

Radio emission of highly inclined cosmic ray air showers measured with LOPES

J. Petrovic^g, W.D. Apel^a, F. Badea^a, L. Bähren^b, K. Bekk^a, A. Bercuci^c, M. Bertaina^d, P.L. Biermann^e, J. Blümer^{a,f}, H. Bozdog^a, I.M. Brancus^c, S. Buitink^g, M. Brüggemann^h, P. Buchholz^h, H. Butcher^b, A. Chiavassa^d, K. Daumiller^a, A.G. de Bruyn^b, C.M. de Vos^b, F. Di Pierro^d, P. Doll^a, R. Engel^a, H. Falcke^{b,e,g}, H. Gemmekeⁱ, P.L. Ghia^j, R. Glasstetter^k, C. Grupen^h, A. Haungs^a, D. Heck^a, J.R. Hörandel^f, A. Horneffer^{e,g}, T. Huege^{a,e}, K.-H. Kampert^k, G.W. Kant^b, U. Klein^l, Y. Kolotaev^h, Y. Koopman^b, O. Krömerⁱ, J. Kuijpers^g, S. Lafebre^g, G. Maier^a, H.J. Mathes^a, H.J. Mayer^a, J. Milke^a, B. Mitrica^c, C. Morello^j, G. Navarra^d, S. Nehls^a, A. Nigl^g, R. Obenland^a, J. Oehlschläger^a, S. Ostapchenko^a, S. Over^h, H.J. Pepping^b, M. Petcu^c, T. Pierog^a, S. Plewnia^a, H. Rebel^a, A. Risse^m, M. Roth^f, H. Schieler^a, G. Schoonderbeek^b, O. Sima^c, M. Stümpert^f, G. Toma^c, G.C. Trinchero^j, H. Ulrich^a, S. Valchierotti^d, J. van Buren^a, W. van Capellen^b, W. Walkowiak^h, A. Weindl^a, S. Wijnholds^b, J. Wochele^a, J. Zabierowski^m, J.A. Zensus^e and D. Zimmermann^h

(a) Institut für Kernphysik, Forschungszentrum Karlsruhe, 76021 Karlsruhe, Germany

(b) ASTRON, 7990 AA Dwingeloo, The Netherlands

(c) National Institute of Physics and Nuclear Engineering, 7690 Bucharest, Romania

(d) Dipartimento di Fisica Generale dell'Università, 10125 Torino, Italy

(e) Max-Planck-Institut für Radioastronomie, 53010 Bonn, Germany

(f) Institut für Experimentelle Kernphysik, Universität Karlsruhe, 76021 Karlsruhe, Germany

(g) Department of Astrophysics, Radboud University Nijmegen, 6525 ED Nijmegen, The Netherlands

(h) Fachbereich Physik, Universität Siegen, 57068 Siegen, Germany

(i) Inst. Prozessdatenverarbeitung und Elektronik, Forschungszentrum Karlsruhe, 76021 Karlsruhe, Germany

(j) Istituto di Fisica dello Spazio Interplanetario, INAF, 10133 Torino, Italy

(k) Fachbereich Physik, Universität Wuppertal, 42097 Wuppertal, Germany

(l) Radioastronomisches Institut der Universität Bonn, 53121 Bonn, Germany

(m) Soltan Institute for Nuclear Studies, 90950 Lodz, Poland

Presenter: S. Buitink (sbuitink@astro.ru.nl), net-buitink-S-abs2-he13-poster

LOPES (LOFAR Prototype Station) is an array of dipole antennas used for detection of radio emission from air showers. It is co-located and triggered by the KASCADE (Karlsruhe Shower Core and Array Detector) experiment, which also provides informations about air shower properties. Even though neither LOPES nor KASCADE are completely optimized for the detection of highly inclined events, a significant number of showers with zenith angle larger than 50° have been detected in the radio domain, and many with very high field strengths. Investigation of inclined showers can give deeper insight into the nature of primary particles that initiate showers and also into the possibility that some of detected showers are triggered by neutrinos. In this paper, we show the example of such an event and present some of the characteristics of highly inclined showers detected by LOPES.

1. Introduction

When a cosmic ray interacts with particles in the Earth atmosphere, it produces a shower of elementary particles propagating towards the grounds with almost the speed of light. The first suggestion that these air showers can produce radio emission was given by Askaryan [1] based on a charge-excess mechanism. Recently, Falcke&Gorham [2] proposed that the mechanism for radio emission of air showers is coherent geosynchrotron

radiation. Secondary electrons and positrons produced in the particle cascade are deflected in the Earth magnetic field and this produces dipole radiation that is relativistically beamed in the forward direction. The shower front emitting the radiation has a thickness which is comparable to a wavelength for radio emission below 100MHz (around few meters). The emission is coherent which amplifies the signal.

Radio emission of cosmic ray air showers has been detected by LOPES (LOFAR Prototype Station) [3], a phased array of dipole antennas co-located with the KASCADE (Karlsruhe Shower Core and Array Detector) experiment which provides coincidence triggers for LOPES and well-calibrated informations about air-shower properties, like electron number N_e , reconstructed muon number N_μ , azimuth and zenith angle of the event. The LOPES experiment and data reduction are described in detail by Horneffer et al. [4].

Highly inclined showers are expected to be very well detectable in the radio domain [5],[6]. However, we have to mention that neither LOPES nor KASCADE are optimized for large zenith angles. For example, KASCADE reconstruction of electron and muon number can be not accurate especially in cases when the shower core of specific event falls out of the KASCADE array.

Inclined cosmic ray air showers are specific, since they travel trough few times longer distances in the Earth atmosphere compared to vertical showers, and due to this most of the electromagnetic (particle) component of those showers has been absorbed. So, inclined showers that start high in the atmosphere (initiated by protons, iron nuclei or gamma-photons) will have large electron deficiency on the ground level compared to vertical showers. On the other hand, neutrino induced showers may be generated at any distance from the ground [7] so they could clearly be distinguished from those whose primary particle initiated the shower high in the atmosphere by the number of electrons that reach the ground level.

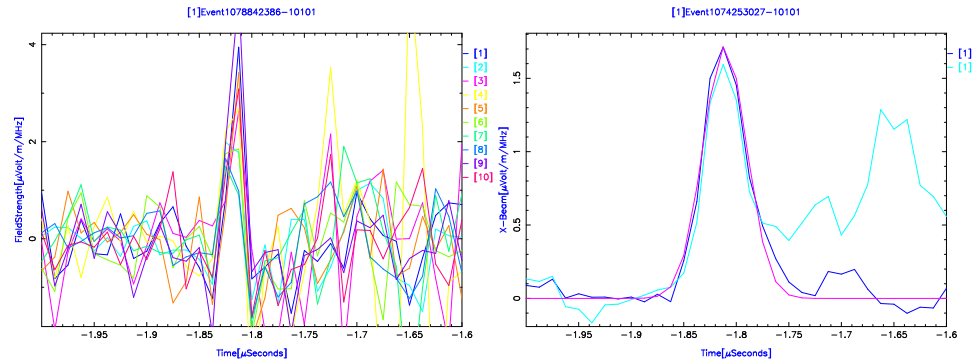


Figure 1. Left: Electric field as a function of time for 10 LOPES antennas for an event detected in March 2004 with zenith angle 53.3° and azimuth 54.7° . The angle between the shower core and the Earth magnetic field is 69.8° . $N_e=1.5 \cdot 10^6$, $N_\mu=1.5 \cdot 10^6$, reconstructed by KASCADE. Right: Radio emission as a function of time after beam forming for the same event. Dark blue line represents X-beam, light blue line the total power and pink line the Gaussian fit for X-beam.

2. Discussion

We made a selection of inclined events from data taken during year 2004 (also detected by the KASCADE array where the reconstructed shower cores were less than 100m away from the array center). We found 2017 events with zenith angle larger than 50° . Then we introduced an additional condition: $N_\mu > 10^5$. In this way we narrowed the selection to 51 event and more than 40% of those are detected in the radio domain, many with very large field strengths, even though a threshold on muon number is lower than the one used for bright

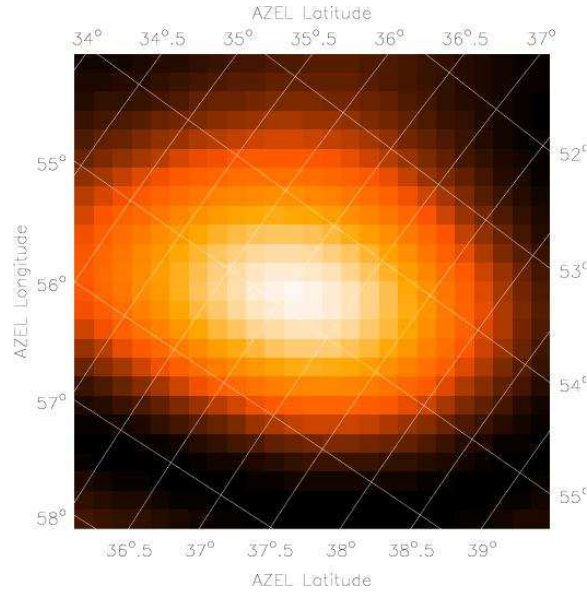


Figure 2. Radio map of the cosmic ray shower with zenith angle 53.3° detected in March 2004. Azimuth (AZEL longitude) and elevation (AZEL latitude) of the event are given on the axes.

events in [3]. However, as we already mentioned, KASCADE is not optimized for large zenith angles, so the reconstruction of the electron and muon number failed for half of the detected events. This leaves us with 10 events with strong radio signal and reliable shower properties reconstructed by KASCADE.

As an example, we show here one of those events, detected in March 2004 with zenith angle 53.3° and azimuth 54.7° with roughly the same number of electrons and muons $N_e, N_\mu \approx 1.5 \cdot 10^6$ (reconstructed by KASCADE). The angle between the shower axis and the Earth magnetic field (geomagnetic angle) is 69.8° .

In Figure 1 (left) we show the electric field as a function of time for each antenna. The field is coherent at $-1.825\mu\text{s}$ which is the arrival time of the shower. The incoherent noise after is radio emission from photomultipliers and in this case is very weak. Figure 1 (right) shows the radio emission as a function of time after X(excess)-beam forming. This beam is formed in the following way. First squared signals of all antennas are summed which gives the total power. Then signals of all two antenna combinations are multiplied and summed which gives the CC(cross correlation)-beam. Finally, CC-beam is multiplied with the ratio between CC-beam and total power. In this way a suppression of incoherent noise is achieved.

Figure 2 is a radio map of the example event. The air shower is the brightest point in the sky for several tens of nanoseconds. The resolution of the map is $\sim 2^\circ$ in azimuth and elevation towards the zenith.

Figure 3 gives the energy deposit of the chosen cosmic ray shower over the KASCADE array with e/γ detectors (left panel) and particle arrival time (right panel). We can see that the shower core falls within the KASCADE array and that the maximum energy is deposited in the north-western part, within the LOPES array. We can notice elliptical shapes of isolines of energy deposit, which is typical for inclined events.

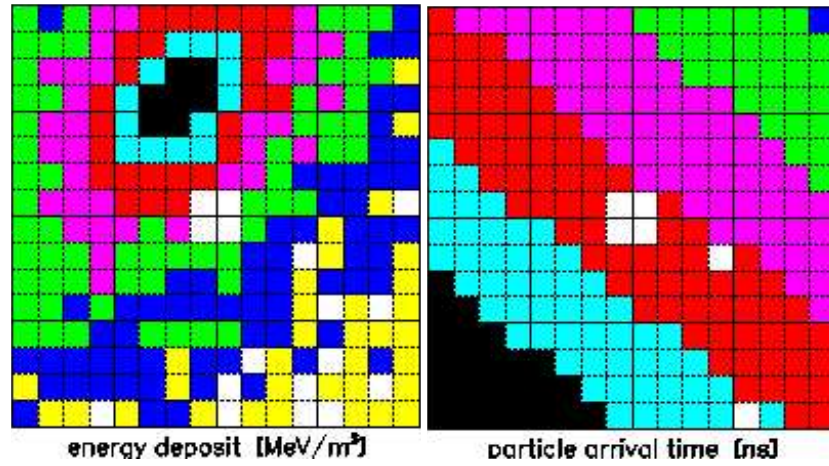


Figure 3. Left: Energy deposit of the cosmic ray shower with zenith angle 53.5° detected in March 2004 for e/γ KASCADE detectors. Dark blue color shows energy deposit of 100MeV/m^2 , red color 1000MeV/m^2 . Maximum energy deposit for this event is $\sim 4500\text{MeV/m}^2$. Right: Particle arrival time for the same event. Dark blue color represents $\sim 300\text{ns}$, red $\sim 800\text{ns}$.

3. Conclusions

Even though neither LOPES nor KASCADE are completely optimized for the detection of highly inclined events, we find that in a selection of events with zenith angle larger than 50° and $N_\mu > 10^5$ (51 event) around 40% of all events is detected in the radio domain, and some of them with very high field strengths, like the example we have presented in this paper. The most inclined cosmic ray air shower that we detected with LOPES has a zenith angle of almost 70° .

After checking the reconstruction of N_e and N_μ different correlations can be considered, for example between radio pulse height and muon number, electron number or geomagnetic angle. This will also give more insight into the nature of primary particles that initiate showers and the possibility that some of detected showers might have been triggered by neutrinos.

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