Figure 5. Lateral distribution of muons obtained by A20 and A100 for vertical showers ($<\sec\theta>$=1.09) in $10^{17.5}$eV $\leq$E$<$ $10^{18.0}$eV (open circles) and $10^{18.5}$eV $\leq$E$<$ $10^{19.0}$eV (closed circles) normalized to $10^{17.75}$eV and $10^{18.75}$eV. The Greisen LDM represented by equation (4) and the modified LDM represented by equation (6) are shown by thick solid lines and thin solid lines respectively. The experimental data are well represented by equation (6).

Figure 6. The average N_μ is plotted as a function of N_e for vertical showers ($<\sec\theta>$=1.05) in (a) and as a function of secθ for inclined showers (1.0 $\leq\sec\theta <$ 1.6) in (b). Data for inclined showers are devided into three size regions (closed circle; 10^7–$10^{7.5}$, open circle; $10^{7.5}$–10^8, open square: 10^8–$10^{8.5}$). Curves drawn are the average relations between N_μ and N_e obtained by A1 and expressed by equation (7).

Figure 7. The average $\rho_\mu(600)$ is plotted as a function of S(600) for A1 (closed circles), A20 (dotted squares), and A100 (open circles) for vertical showers ($<\sec\theta>$=1.09). Data from A100 are normalized to those from A20. Solid lines drawn represent equation (8).

Figure 8. LDM obtained by A20 for vertical showers ($<\sec\theta>$=1.09) is compared with those obtained by the Yakutsk group and the Sydney group . Data from A20 are plotted with open circles normalized to $10^{18.75}$eV and data from the Yakutsk group are represented by a dotted line. The solid line is the average LDM of the Sydney group.

Figure 9. The average ratios of $\rho_\mu(600)$ to S600 obtained by A1 (closed circle), A20 (dotted squares) and A100 (open circle) are plotted as a function of S(600) in five different zenith angle bins between 1.0 $\leq\sec\theta <$ 2.0. A solid line in each diagram is the best fit to the experimental data in each zenith angle bin. Values of S(600) corresponding to 5×10^{18}eV are shown by arrows for each zenith angle. The result from the Yakutsk group is shown by a dotted line in (e).

Figure 10. Lateral distributions of muons and charged particles obtained by the Akeno experiment are compared at the atmospheric depth of 1000gcm^{-2} with the simulation results of protons at 10^{19}eV by Cronin using the "MOCCA" program provided by Hillas. A closed circle shows the avereage energy deposit per square meter in a 50mm thick scintillator and an open circle corresponds to the muon density for >1GeV. Lateral distributions of charged particles and muons obtained by this experiment are shown by a thin solid line and a thick solid line respectively.

Figure 11. $\rho_\mu(600)$'s divided by the expected values are plotted as a function of S(600) between $10^{16.5}$eV and 10^{19}eV. The expected densities are extrapolations from the average relation determined below $10^{17.5}$eV. Data from A1 (closed circles), A20 (open circles) are included. Data from A100 are combined with those from A20 after normalizing the threshold energy of muons. Values of S(600) corresponding to $10^{17.5}$eV and $10^{18.5}$eV are shown by arrows. An expectation from the composition model proposed by the Fly's Eye group is shown by dotted lines (90% C.L.), chain lines (68% C.L.) and a solid line (the average) respectively.

Figure captions

Figure 1. A schematic diagram of the Akeno Giant Air Shower Array (AGASA, A100). The relationship among AGASA (A100), the 1km^2 array (A1), and the 20km^2 array (A20) is shown.

Figure 2. The average lateral distribution of muons obtained by A1 for vertical air showers (<secθ >=1.05). Data are divided into eight logarithmically equal size bins between 10^7 and 10^9 (cross; 10^7-$10^{7.25}$, diamond; $10^{7.25}$-$10^{7.5}$, closed circle; $10^{7.5}$-$10^{7.75}$, open circle; $10^{7.75}$-$10^{8.0}$, closed square; $10^{8.0}$-$10^{8.25}$, open square; $10^{8.25}$-$10^{8.5}$, closed triangle; $10^{8.5}$-$10^{8.75}$, open triangle; $10^{8.75}$-$10^{9.0}$). The best fit curves are shown by solid lines in each size bin.

Figure 3. Zenith angle dependence of the lateral distribution of muons for 1.0 $\leq$secθ < 1.6 and $10^{7.0}$ $\leq$N$_e$ < $10^{7.5}$ in (a) and $10^{7.5}$ $\leq$N$_e$ < $10^{8.0}$ in (b). The best fit curve is shown in each zenith angle bin (cross; 1.0$\leq$secθ <1.1, diamond; 1.1$\leq$secθ <1.2, closed circle; 1.2$\leq$secθ <1.3, open circle; 1.3$\leq$secθ <1.4, closed square; 1.4$\leq$secθ <1.5, open square; 1.5$\leq$secθ <1.6).

Figure 4. An example of the lateral distributions of charged particles (closed circles) and muons (open circles) are shown for one of the biggest showers observed by A1. Shower parameters determined are log(E)=18.79 and secθ=1.06.

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A1 data.

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a 90% C.L. in each S(600) are determined and shown in Figure 11 by chain and dotted lines repectively. The deficit ranges estimated by the simulation are listed in Table 5 at 68% C.L.. The deficit in the observed to expected ratio ranges from from −4.8% to −1.2% for E$\geq 10^{17.5}$eV at 68% C.L. over the 50 simulations. The experimental result is consistent with this range. However the deficit extends from −25.8% to −15.0% for E$\geq 10^{18.5}$eV at 68% C.L.. The range corresponding to a 90% C.L. is −31.2% to −11.1% while the experimental deficit is significantly smaller than this value. Therefore the simple two component model proposed by the Fly's Eye group can not be supported by our experiment under the assumptions of the present simulation.

The relation between $\rho_\mu(600)$ and S(600) obtained from the present experiment is expressed by equation (8) and the exponent (0.82±0.03) is in a good agreement with the simulated one, 0.84. If we use assumption 6 of the simulation above and relation (17), the expected ratio for a changing composition given by $\frac{\delta \ln A}{\delta \ln E} = -0.58$ almost coincides with the average line of Figure 11 and is consistent with the simulation. In order to satisfy our experiments, $\frac{\delta \ln A}{\delta \ln E}$ may be more than -0.3 which is near to an upper chain line in Figure 11, that is, a rate of change of composition may be less than half of the Fly's Eye interpretation. We need further accumulation of data and more detailed simulations in order to determine the composition of primary cosmic rays by measuring relations between $\rho_\mu(600)$ and S(600).

5. Summary

The LDM was determined first by experiment A1 for core distances ≤800m and $10^{16.5}$eV ≤E< $10^{18.5}$eV for both vertical and inclined showers. The parameters which determine the LDM are consistent with those reported previously and are expressed by equation (4) (Hayashida *et al.*, 1991). However the LDM for core distances beyond 800m becomes significantly steeper compared with the extrapolation of equation (4) for showers between $10^{17.5}$eV and $10^{19.5}$eV. The shape of the LDM is in agreement with both the Yakutsk result and the simulation results from the MOCCA program for core distances beyond 300m within statistical errors. The absolute value is, however, larger than both of them by about a factor of 2.0∼2.5.

The energy dependence of the ratio of $\rho_\mu(600)$ to S(600) is examined for five zenith angle ranges and no significant deviation from the expected dependence is observed up to 10^{19}eV for sec$\theta < 2.0$.

The relation between N_μ and N_e is expressed in equation (7) for $10^{16.5}$eV ≤E< $10^{18.0}$eV where N_μ is found to be proportional to $N_e^{0.77}$ for vertical showers. At higher energies it is shown that $\rho_\mu(600)$ is proportional to $S(600)^{0.82}$ for $10^{17.5}$eV ∼ $10^{19.0}$eV. The systematic change in the chemical composition of cosmic rays claimed by the Fly's Eye group above $10^{17.5}$eV under the assumption of a simple two component model was investigated by using the relation between $\rho_\mu(600)$ and S(600). This change in the composition can not be confirmed within the present experimental uncertainties. The present experiment suggests a much smaller rate of change of composition between $10^{17.5}$eV and $10^{18.5}$eV.

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Table 5. $\rho_\mu(600)$ measured by each detector in A1 is summed up and compared with the expected value. "Deficit" is defined in the text. The deficit ranges were simulated under the assumption of the two component composition model proposed by the Fly's Eye group and are listed at 68% confidence level.

Experiment	$\sum_{total\ events} \rho_\mu(600)$	Deficit	Simulation
A1 ($>10^{17.5}$eV)	3.95×10^4	(−1.5∼−0.8)±0.5%	(−4.8∼ −1.2)%
($>10^{18.5}$eV)	2.45×10^3	(−2.8∼ 0.4)±2.0%	(−25.8∼−15.0)%

Under the assumption of the two component composition model proposed by the Fly's Eye group, the expected $\rho_\mu(600)$ vs S(600) relation in the present experiment was estimated by a Monte Carlo simulation with the following assumptions.

1.	$S(600)_{Fe}/S(600)_p = 1.04$;	from Cronin (1994b), Dawson (1993)
2.	Fluctuation of $S(600)_{Fe}$ = 15%; $S(600)_p$ = 15%	from Figure 2(a) in Dawson (1993)
3.	$\rho_\mu(600)_{Fe}/\rho_\mu(600)_p = 1.7$;	from Cronin (1994b), Dawson (1994)
4.	Fluctuation of $\rho_\mu(600)_{Fe}$ = 10%; $\rho_\mu(600)_p$ = 15%	from Figure 2(a) in Dawson (1993)
5.	$S(600)_p$ and $S(600)_{Fe}$ are proportional to $E^{1.0}$;	from Figure 3 in Dawson (1993)
6.	$\rho_\mu(600)_p$ and $\rho_\mu(600)_{Fe}$ are proportional to E^α, $\alpha = 0.84$;	from Dawson (1994)
7.	Energy dependence of the composition of primary cosmic rays;	a two composition model proposed from the Fly's Eye experiment (Bird *et al.*, 1994)
8.	Absolute value of $\rho_\mu(600)$;	normalized to the experiment at $10^{17.5}$eV
9.	Primary energy spectrum;	Akeno spectrum from Yoshida *et al.* (1995)
10.	The error in determining of S(600); ($\sigma_{S(600)}$)	20% (Section 2.1) S(600) is derived from N_e in case of the A1.
11.	The determination of $\rho_\mu(600)$;	determined from experimental points from 25m^2 detectors between 500m and 800m from the core by interpolating with LDM expressed by equation (6).
12.	The error in determining of $\rho_\mu(R)$;	Poisson distribution over 25m^2 area. On-Off density is used for vertical showers.

In this simulation, events were accumulated until the total number of events amounted to the same number as seen by the experiment. The relation between $\rho_\mu(600)$ and S(600) was determined below $10^{17.5}$eV in each simulation and the $\rho_\mu(600)$ values from more energetic showers were compared with those expected from the extrapolation of this relation. In total 50 data-sets were simulated. A 68% confidence level (C.L.) and

or,

$$X_{max} = a \cdot ln\left(\frac{\mathrm{E}}{\mathrm{A}}\right) + b.$$

Similarly,

$$\frac{\delta \mathrm{lnN}_\mu}{\delta \mathrm{lnE}} = \alpha + (1-\alpha)\frac{\delta \mathrm{lnA}}{\delta \mathrm{lnE}} \tag{15}$$

where α is related by

$$\mathrm{N}_\mu = k(A)\mathrm{E}^\alpha.$$

Therefore, if the change of elongation rate around 3×10^{17}eV observed with the Fly's Eye data is due to a change in composition ($\frac{\delta \mathrm{lnA}}{\delta \mathrm{lnE}} \sim -0.58$), a change of slope in the N_μ vs E relation may be expected around 3×10^{17}eV. In the following, any changes in $\frac{\delta \mathrm{lnN}_\mu}{\delta lnE}$ are investigated by replacing N_μ by $\rho_\mu(600)$ and E by S(600).

Relations between the densitites of muons (ρ_μ) and those of all charged particles (S) at various core distances were shown by Dawson (1993) for simulated proton and iron primaries. There is a factor of 1.7 difference in ρ_μ for iron and proton primaries at constant S, suggesting the view that there must be an observable change in the slope of the N_μ vs S(600) relation, if there is a change of composition from a predominatly heavy to a light composition. If there were a systematic change in the chemical composition of primary cosmic rays between $10^{17.5}$eV and $10^{18.5}$eV, a possible change in the relation between N_μ and N_e (equation (7)), or in the relation between $\rho_\mu(600)$ and S(600) (equation (8)) may be expected.

The relation between $\rho_\mu(600)$ and S(600) was examined for vertical showers ($\sec\theta < 1.2$) and was compared with the extrapolated one from the relation determined in the energy region below $10^{17.5}$eV. Expected values of $\rho_\mu(600)$ were derived from the results of A1 for $-0.5 \leq \log \mathrm{S}(600) < 0.2$. Contributions from accidental backgrounds exceed 5% of densities measured for $\log \mathrm{S}(600) < -0.5$ and hence data in this energy region were not used. $\rho_\mu(600)$'s for A1, A20, and A100 normalized to those of A20, were divided by the expected muon density given by equation (8) with b=0.82 over whole S(600) range and are plotted in Figure 11. Data from A100 are combined with those from A20. Two vertical arrows in the figure correspond to $10^{17.5}$eV and $10^{18.5}$eV respectively. No systematic change in the ratio of the observed to expected $\rho_\mu(600)$'s can be seen for $\mathrm{E} \geq 10^{17.5}$eV.

Figure 11

In order to examine a possible change quantitatively, $\rho_\mu(600)$ measured by each detector in A1 is summed up for all the events with $\mathrm{E} \geq 10^{18.5}$eV and is 2449. The value expected from the extrapolation of the relation between $\rho_\mu(600)$ and S(600) determined in the energy range below $10^{17.5}$eV is estimated to be 2440$\sim$2517. The decreased proportion of observed muon particles compared with the expected number (hereafter we call as "deficit") is $(-2.8 \sim 0.4)\% \pm 2.0\%$. The number of muons above $10^{17.5}$eV is 3.949×10^4, whereas $(3.983\sim4.007)\times10^4$ are expected from the linear extrapolation of the lower energy results. The observed number shows a deficit of $(-1.5 \sim -0.8)\% \pm 0.5\%$. These comparisons of integrated $\rho_\mu(600)$'s with the expected ones are presented in Table 5.

4.4. Comparison of the present LDM with the simulated LDM

The LDM beyond 800m from the core can not be expressed by Greisen's formula (equation (4)). It becomes significantly steeper than the extrapolated values of equation (4) and the muon density at R=2000m is about 20% of the extrapolated value. Simulated values at distances far from the core for EAS of energy 10^{19}eV are available from simulations made by Cronin (1994a) at the atmospheric depth of 1000gcm^{-2} using a program prepared by Hillas (MOCCA).

The lateral distributions of both charged particles and muons are compared simultaneously in form of Density$\times$R^2. The lateral distribution of charged particles obtained by AGASA has been reported in Yoshida *et al.* (1994). S(600) at 1000gcm^{-2} is evaluated as 36.5m^{-2} at 10^{19}eV with the difference of Moliere length taken into account. $\rho_\mu(600)$ at 1000gcm^{-2} is derived from equation (8) using 42.5 as S(600), which is obtained at secθ=1.09 at Akeno level for 10^{19}eV. The Akeno results are shown in Figure 10 by a thin solid line for charged particles and by a thick solid line for muons. The average energy deposit per square meter in a 50mm thick scintillator calculated by the MOCCA program is compared with the lateral distribution of charged particles obtained by Akeno by a filled circle. The simulated muon density (>1GeV) is also shown in the same figure as an open circle.

Figure 10

The charged particle density obtained by the Akeno result is higher than the prediction of MOCCA by a factor of 1.4 at 600m from the core and 1.9 at 2000m. The Akeno result for muon densities is higher than that predicted by MOCCA by a factor of 2.0 and 2.3 at 600m and 2000m from the core respectively. If the density of charged particles is normalized at 600m from the core, the muon density obtained by the Akeno experiment is higher by a factor of 1.5 than the prediction of MOCCA at 600m from the core. Except for these differences of absolute values, the lateral distributions of both charged particles and muons obtained by Akeno are rather consistent with the prediction by MOCCA simulations in shapes.

In Table 4 are listed typical densities (per m^{-2}) at secθ=1.0 estimated from the present experiment.

Table 4. Typical densities (per m^{-2}) at secθ=1.0 estimated from the present experiment.

Primary energy	S(600)	S(1000)	S(1500)	$\rho_\mu(600)$	$\rho_\mu(1000)$	$\rho_\mu(1500)$
1×10^{19}eV	50	6.6	1.1	3.4	0.57	0.10
3×10^{19}eV	150	20	3.4	8.4	1.4	0.25
1×10^{20}eV	500	66	11	22	3.7	0.66

4.5. Composition of cosmic rays

Assuming the validity of the superposition model for the development of showers initiated by nuclei, the elongation rate ($\frac{\delta X_{max}}{\delta \ln E}$) can be related to a mass A by

$$\frac{\delta X_{max}}{\delta \ln E} = a\left(1 - \frac{\delta \ln A}{\delta \ln E}\right),$$

4.3. Conversion between N_μ and E

In order to determine the energy spectrum of primary particles from the N_μ spectrum (as done in the Sydney experiment), the conversion relation from N_μ to E or from $\rho_\mu(600)$ to E is necessary. In section 3.3, the conversion relation from N_e to N_μ was obtained by A1 as equation (7), while N_e is converted into E using equation (1). These two equations are combined into

$$\log N_\mu = (0.84 \pm 0.02)\log E - (8.38 \pm 0.35) \tag{11}$$

at secθ=1.0. The conversion relation from N_μ to E is expressed by

$$E = 2.16 \times 10^{18}(N_\mu/10^7)^{1.19} eV \tag{12}$$

This relation is in agreement with our previous result within experimental errors (Nagano *et al.*, 1984b).

The conversion relation from N_μ to E derived by the Yakutsk group is expressed as

$$E = 2.48 \times 10^{18}(N_\mu/10^7)^{1.05} eV.$$

After taking into account the difference in the depths of the Yakutsk and Akeno experiments (100g/cm^2), we find good agreement in the energy conversion relations, even though the LDM's are inconsistent. Good agreement is not found with the relation used by the SUGAR experiment, which is derived from simulations (model E) by Hillas (1971),

$$E = 1.64 \times 10^{18}(N_\mu/10^7)^{1.075} eV.$$

The relation between $\rho_\mu(600)$ and S(600) obtained by A20 and A100 can be converted into one between E and ρ_μ by using equation (2) and equation (8). The relation at secθ=1.0 is

$$E = 2.22 \times 10^{18}\rho_\mu(600)^{1.22} eV \tag{13}$$

Assuming β, R_0, and δ in equation (6) are independent of primary energy, this relation can be converted to a relation between N_μ and E at secθ=1.0,

$$\log N_\mu = (0.82 \pm 0.03)\log E - (8.06 \pm 0.04) \tag{14}$$

The relations (11) and (14) are smoothly connected with each other within experimental errors for the range $10^{16.5}$eV $\leq$E$<$ $10^{19.5}$eV, although the overlapping region is only $10^{17.5}$eV $\leq$E$<$ $10^{18.5}$eV.

Figure 8 by a dotted line. The modification factor is $(1+R/2000)^{-1}$ instead of $(1+(R/800)^3)^{-0.6}$ and β is a function of secθ instead of R_0 . The shape of the LDM is consistent with the results of A20 over the core distances concerned. However the Yakutsk LDM extrapolated inside 300m from the core is much steeper than the Akeno result. The absolute value of the LDM obtained by the Sydney group is smaller than that of A20 by a factor of 1.5 and that of the Yakutsk is smaller by a factor of 1.8, if compared at R = 600m from the core.

On the other hand, N_μ at $10^{18.75}$eV obtained by the Akeno, Yakutsk, and Sydney groups using their LDM's are 2.5×10^7, 2.2×10^7, and 3.1×10^7 respectively. These values do not differ much in this energy region, even though the shapes of the LDM are different and the primary energy assignment method is different in each experiment. The conversion relation between N_μ and E is discussed in Section 4.3.

4.2. Comparison of ρ_μ(600)/S(600) with the Yakutsk experiment

In order to examine changes of shower characteristics in the energy region above 5×10^{18} eV claimed by Glushkov *et al.*, the average ratios of $\rho_\mu(600)$ to S(600) obtained by A1, A20, and A100 are plotted as a function of S(600) in Figure 9 for five different zenith angle bins. S(600) values corresponding to 5×10^{18}eV at each zenith angle bin are shown by arrows in the figures. In order to convert S(600) to a primary energy, an attenuation of S(600) described in Yoshida *et al.* (1994) is taken into account. No significant change of slopes of the ratios is observed even above 5×10^{18}eV. The ratios are fitted to the following equation and are shown with solid lines.

$$\log\rho_\mu(600)/S(600) = A\log S(600) + B \tag{10}$$

Figure 9

Values of A and B derived from a least square method are listed in Table 3.

The ratio from the Yakutsk experiment for $1.8 \leq \sec\theta < 2.0$ is estimated from Figure 10 of Glushkov *et al.* (1994) and is drawn by a dotted line in Figure 9(e). Our present results for E$\geq 5\times10^{18}$eV are consistent with extrapolations from lower energy regions for secθ <2.0. The behavior seen in the Yakutsk results has not been observed. The confirmation of the present result, however, needs further accumulation of data.

Table 3. Values of A and B defined in equation (10), which show the relation between $\rho_\mu(600)/S(600)$ and S(600), are listed as a function of secθ.

secθ	A	B
1.0 – 1.2	−0.17 ± 0.01	−0.76 ± 0.01
1.2 – 1.4	−0.16 ± 0.01	−0.58 ± 0.01
1.4 – 1.6	−0.17 ± 0.04	−0.40 ± 0.02
1.6 – 1.8	−0.22 ± 0.10	−0.30 ± 0.03
1.8 – 2.0	−0.30 ± 0.22	−0.21 ± 0.06

A1. If the LDM does not change in the energy region concerned, the muon density at 600m from the core, $\rho_\mu(600)$ is expected to be an estimator of N_μ, in the same way as S(600) is an estimator of E_p. $\rho_\mu(600)$ is derived from an observed density at 500∼800m from the core, which is normalized to that at 600m by using equation (6).

Figure 7 shows the relation between $\rho_\mu(600)$ and S(600) obtained by both A20 and A100 for vertical showers ($\sec\theta < 1.2$) for $10^{17.5}$eV $\leq$E$< 10^{19.5}$eV. Densities from A100 are normalized to those from A20 ($\rho_\mu(600)$=2.66) and the slopes of the $\rho_\mu(600)$ versus S(600) relations from A100 and A20 agree with each other. The average $\rho_\mu(600)$ at $<\sec\theta>$=1.09 is expressed as a function of S(600) by

$$\rho_\mu(600) = (0.16 \pm 0.01) \times \mathrm{S}(600)^{0.82\pm0.03} \tag{8}$$

The $\rho_\mu(600)$ data obtained from A1 for the same zenith angle bin are also shown in Figure 7, and are consistent with the other plots. The average $\rho_\mu(600)$/S(600)'s at $<\sec\theta>$=1.09 for A20 and A100 are expressed as

$$\rho_\mu(600)/\mathrm{S}(600) = (0.17 \pm 0.01) \times \mathrm{S}(600)^{-0.20\pm0.02} \tag{9}$$

Figure 7

4. Discussion

4.1. Comparison of LDM with other experiments

In Figure 8, $\rho_\mu \times \mathrm{R}^2$ for vertical showers ($\sec\theta$ <1.2) obtained by A20 normalized to $10^{18.75}$eV is plotted to display the distribution more clearly. The vertical atmospheric depth of Akeno is 920gcm^{-2} and is shallower than that of Sydney (980gcm^{-2}) and Yakutsk (1020gcm^{-2}). In the figure, results of these two experiments at the atmospheric depth of 1020gcm^{-2} are compared, where the threshold energy of muons in the Sydney experiment and that of the Yakutsk experiment are 0.75secθ GeV and 1.0 secθGeV respectively. The corresponding secθ's at the atmospheric depth of 1020gcm^{-2} are 1.11 for Akeno, 1.04 for Sydney, and 1.0 for Yakutsk. The average zenith angle of the Akeno results is $\sec\theta$ = 1.09 and does not differ much from secθ=1.11.

Figure 8

The LDM of the Sydney group is expressed by Greisen's formula in which β in equation (4) is a function of secθ with a constant R_0 (Winn *et al.*, 1986) and is shown in the same figure by a solid line. The data from A20 show a steeper lateral distribution than the Sydney LDM. In the case of the Sydney experiment, there may be a triggering bias favouring showers with flatter muon lateral distributions. Furthermore, since only muon detectors are used in that experiment, the determination of the core position and the LDM are correlated. On the other hand the LDM from A20 is determined independently of the array triggering and core position determination.

The LDM of the Yakutsk group recently reported (Glushkov *et al.*, 1994) is a modified version of Greisen's formula similar to equation (6) and is shown in

inside A1, whose positions are 1.5∼2.5km distant from the center of A20. Large area muon detectors are advantageous in investigating the LDM with R exceeding ∼1000m.

The average ρ_μ's obtained from A20 are plotted with open and closed circles as a function of core distance in Figure 5 for two energy ranges. Here the ρ_μ's are normalized to energies of $10^{17.75}$eV and $10^{18.75}$eV respectively and the average values are plotted. The thick solid lines are those expected from equation (4) using β and R_0 obtained for R≤800m and it can be seen the data agree very well with these lines within 800m from the core. They show, however, significantly lower values than the extrapolated values of the equation (4) for R>800m. The best fitted LDM out to 2000m, shown by thin lines, is expressed by modifying equation (4) to give,

$$\rho_\mu = \mathrm{N}_\mu(C'_\mu/R_0^2)\mathrm{r}^{-0.75}(1+\mathrm{r})^{-\beta}\{1+(\mathrm{R}/800m)^3\}^{-\delta} \tag{6}$$

where $\delta \sim 0.6$. C'_μ was numerically calculated to be 0.325 for β=2.52. Independent measurements were also made for vertical showers using A100. In order to take into account the difference in muon threshold energies of A20 (>1GeV) and A100 (>0.5GeV), the densities from A100 were reduced by 1.4, a value obtained from a previous experiment (Matsubara *et al.*, 1985). Muon densities from A100, adjusted to a threshold of 1GeV, are plotted in Figure 5 with squares for $10^{18.5}$eV∼10^{19}eV, and are consistent with ρ_μ's from A20 within experimental errors.

Figure 5

The number of inclined showers (secθ ≥1.2) is not large enough to determine the zenith angle dependence of the LDM out to 2000m. The zenith angle dependece is discussed in terms of $\rho_\mu(600)$ in Section 4.2.

3.3. The relation between N_μ and N_e

The average N_μ is determined for data obtained by A1 by integrating equation (4) for a number of N_e bins using β and R_0 obtained above. The average N_μ is plotted in Figure 6(a) as a function of N_e for vertical showers (secθ <1.1). N_μ is also plotted in Figure 6(b) as a function of secθ for three different N_e bins. N_μ is expressed as a function of N_e and secθ by

$$\mathrm{N}_\mu = (2.6 \pm 1.4) \times 10^{5.0+\mathrm{a}}(\mathrm{N}_e/10^7)^{\mathrm{b}} \tag{7}$$

for $10^{7.0} \leq \mathrm{N}_e < 10^{8.5}$ and secθ ≤1.6, where a = (1.07±0.13)(secθ−1) and b = (0.77±0.02)− (0.17±0.02)(secθ−1).

Figure 6

3.4. The relation between $\rho_\mu(600)$ and S(600)

In the case of A20, the 1km separation of muon detectors makes it difficult to integrate equation (6). It is also difficult to determine N_μ with A100 (for E≤$10^{19.0}$eV) since the detectors have much smaller area and are separated by larger distances compared with

3. Results

3.1. The lateral distribution of muons (LDM) for core distances $\leq$800m

First the LDM for core distances(R) within 800m from the core was studied. For this purpose a data set obtained by A1 was used. In Figure 2, the average ρ_μ's in $\Delta\log N_e$=0.25 bins from $\log N_e$=7.0 to $\log N_e$=9.0 are plotted for vertical showers as a function of core distance. In each bin ρ_μ's are normalized to the middle value of N_e. Those in $\Delta\sec\theta$=0.1 bins are shown in Figure 3(a) and (b) for two N_e ranges.

Figure 2, Figure 3

To examine the N_e dependence or zenith angle dependence of the shape of the LDM, the formula given by Greisen (1960) was used,

$$\rho_\mu(\mathrm{R}) = \mathrm{N}_\mu(C_\mu/R_0^2)\mathrm{r}^{-\alpha}(1+\mathrm{r})^{-\beta} \tag{4}$$

where N_μ is the total number of muons, r=R/R_0 and R_0 is a characteristic distance. C_μ is a normalization factor and is given by $C_\mu = \Gamma(\beta)/\{2\pi\Gamma(2\text{–}\alpha)\ \Gamma(\alpha+\beta\text{–}2)\}$. In the present analysis, the value of α was fixed to be 0.75 as determined by a previous analysis (Hayashida *et al.*, 1991).

β and R_0 were determined for vertical showers ($\sec\theta < 1.1$) by fitting ρ_μ to equation (4) using a least square method in each N_e bin. The best fitted values are β=2.52$\pm$0.04 and R_0=266$\pm$32m. The lines in Figure 2 are drawn for the best fitted values of β and R_0 in each bin and no significant energy dependence has been found for either β or R_0. The zenith angle dependence of R_0 was determined using inclined showers with $\sec\theta$ <1.8 by keeping β=2.52 constant for N_e between $10^{7.0}$ and $10^{9.0}$, and is expressed as

$$\log(R_0) = (0.58 \pm 0.04)(\sec\theta - 1) + (2.39 \pm 0.05) \tag{5}$$

where C_μ in the equation (4) is 0.262. In general, increasing (or decreasing) β in equation (4) gives rise to a similar effect on the LDM as decreasing (or increasing) R_0. R_0 may be understood as a parameter reflecting the muon production height. In Figure 3(a) and (b) the curves are drawn using this equation. It is found that the LDM derived by Greisen in a much lower energy region is still a good expression for the showers of the present energy region within 800m from the core. ρ_μ can not be determined in this experiment for R> 800m owing to the restricted size of the 1km^2 air shower array.

An example of a large shower observed by A1 is shown in Figure 4.

Figure 4

3.2. LDM for core distances >800m

The LDM for core distances 250m $\leq$R< 2,000m was investigated using data obtained by A20 for $10^{17.5}$eV $\leq$E< $10^{19.5}$eV for vertical showers ($\sec\theta < 1.2$). In the case of A20, the average ρ_μ's are determined with 8 muon detectors of 25m^2 area deployed

Table 2. The number and the are of muon detectors used for the present analysys.

Array	Area of one detector (m^2)	Number of detectors	Length of each PC (m)	Number of PC in each detector	Threshold muon energy (GeV)
A1	25.0	8	5.0	50	1.0
A20	25.0	8	5.0	50	1.0
A100	2.8	12	2.0	14	0.5
	10.0	2	5.0	20	0.5

tenths of a particle equivalent. The pulse width is proportional to the logarithm of the number of particles incident on a counter. The dynamic range of each counter is about $\sim$100 particles. The pulse width of each counter is recorded when trigger conditions of air showers are satisfied in A1 and A20. In the case of A100, data are recorded whenever more than two counters separated by at least 10cm are hit, or when the surface detector placed at the same position is hit. The gate widths for muon data taking are 8μsec for A1, 64μsec for A20, and 15.5μsec for A100. A short gate width can be used in A100 because the trigger time is assigned after adjusting for the delay time due to the cable length. This time is then examined in the stored data in cyclic memory. Accidental backgrounds were examined experimentally using random triggers for the three experiments separately and were determined to be $0.003\mathrm{m}^{-2}$ for A1, $0.022\mathrm{m}^{-2}$ for A20, and 0.003 or $0.006\mathrm{m}^{-2}$ for A100, depending on materials used for the absorber (a combination of iron and lead, or concrete shielding.)

The muon density (ρ_μ) in each detector is calculated by two methods, the so-called On-Off Density and Analog Density method. The On-Off Density is determined from the number of hit counters under an assumption that the number of particles incident on each counter follows a Poissonian distribution around the average value. In the case of n hit counters out of m available counters, ρ_μ is given by

$$\rho_\mu = -m\ln(1-n/m)/Area \tag{3}$$

where $Area$ is the total area of the detector. As long as showers are vertically incident and n is much smaller than m, this On-Off Density is a good measure of a real density. On the other hand, the Analog Density is determined by the sum of energy losses in all counters divided by the average energy loss of a vertically traversing muon. Since the pulse height distribution is asymmetrical with a long tail, reflecting a Landau fluctuation of energy losses in gas, the Analog Density gives a larger value than the On-Off Density. This systematic factor depends on the distance from the core and is derived by comparing the Analog Density with the On-Off Density. At distances more than a few 100m from the core, this value is 1.6. In this paper, the On-Off Density is used only for vertical showers and the Analog Density is used for both vertical and inclined showers.

In the present analysis, the lateral distribution of muons was not determined for each shower. Muon densities, ρ_μ, are accumulated in a number of bins representing ranges of zenith angles (θ), N_e's, local densities at 600m from the core (S(600)), and core distances (R). After averaging ρ_μ in each bin, the average backgrounds are subtracted from the average densities.

since June 1991. The absorber is either concrete shielding or a combination of iron and lead, and the threshold energy is about 0.5GeV. Proportional counters of the same type found in A1 are used in AGASA and the length of each counter is either 2m or 5m, and the area of each detector is $2.8m^2$, or $10m^2$. The number of muon detectors was increased to 27 in April, 1993, but the added detectors are not included in this analysis. The whole array is divided into four branches and data are recorded whenever 5 or more neighboring detectors are triggered within each branch. The number of events analysed is 14735 up to April 1993.

In all three experiments primary energies, core positions, and arrival directions of the showers are determined by the surface detectors and only showers with their core well inside each array are selected for further analysis. For A1 showers, N_e is used as an estimator of primary energy (E). The relation between N_e and E is derived in a previous experiment (Nagano *et al.*, 1984b) and is given by

$$E = 3.9 \times 10^{15} (\frac{N_e}{10^6})^{0.9} \quad eV \tag{1}$$

For showers measured by A20 and A100, the surface detector separations are about 1km and hence N_e can not be determined in individual showers. Therefore the charged particle density at 600m from the core, S(600), which depends rather weakly on primary composition and the stage of shower development (Hillas *et al.*, 1971 and Dai *et al.*, 1988), is used to estimate E:

$$E = 2.0 \times 10^{17} S(600)^{1.0} \quad eV \tag{2}$$

In Table 1 are summarized the durations and the number of events analysed in each array.

Table 1. The durations and the number of events analysed in each array.

Array	Duration of the experiment	Number of Events	Energy Range
A1	October 1981 ~ July 1992	995 457	3×10^{16} ~3×10^{18} eV
A20	December 1984 ~ July 1992	43482	3×10^{17} ~3×10^{19} eV
A100	June 1991 ~ April 1993	14735	3×10^{17} ~3×10^{19} eV

2.2. Muon Detector

The number and the area of muon detectors used for the present analysis are listed in Table 2 together with their threshold energy.

The gain of each proportional counter used in these three experiments is adjusted to within ±3% by measuring the pulse height distribution of omnidirectionally traversing muons through the counter. It has been confirmed experimentally that the peak position for omnidirectionally traversing muons is the same as that for vertically traversing muons. The average pressure of P10 gas in each counter is set at a few % larger than the atmospheric pressure at the Akeno altitude (920 g/cm^2) at around 20°C. The output signal of each counter is shaped to an exponential form with a time constant of 20μsec and discriminated to give a square pulse at the level of a few

$\leq$800m. The relation between N_μ and N_e and between $\rho_\mu(600)$ and S(600) are derived using the lateral distribution thus obtained, where $\rho_\mu(600)$ and S(600) are densities of muon and charged particles on the surface at 600m from the core, respectively. In Section 4, the present results are compared with those given by other experiments and with simulated results. The relation between $\rho_\mu(600)$ and S(600) is discussed in the light of primary cosmic ray composition at the end of the section.

2. Experiment

2.1. General

The experiment has been done at Akeno at 900m above sea level. The whole array consists of several arrays of different detector separations in order to cover the wide primary energy range. The arrays related to the present experiment are shown in Figure 1.

Figure 1

The 1km^2 air shower array (hereafter denoted as A1, Hara *et al.* 1979) consists of 156 scintillation counters of 1m^2 area each deployed on the surface over about 1km^2 area and 8 muon detectors of 25m^2 area each. The surface detectors have been operated since 1979 and the muon detectors since October 1981. The triggering requirement for the present analysis is a coincidence of at least seven out of 37 surface detectors mutually separated by 120m. This array is designed for studies of EAS between $10^{16.5}$eV and $\sim 10^{18.5}$eV. The absorber of the electromagnetic component over the muon detectors is concrete shielding of 2m thickness and the threshold energy of muons is 1GeV. Each muon detector consists of 50 rectangular proportional counters with an area of 5m$\times$10cm and 10cm thickness containing P10 gas (Ar; 90%, CH_4; 10%). The detailed characteristics of this type of proportional counter (PC) are described in Hayashida and Kifune (1980). The lateral distribution of electrons is determined with 156 surface detectors and the uncertainty of N_e determination is 20%. The lateral distribution of muons is determined between 50m and 800m from the core and the uncertainly of N_μ determination of each shower is 20% so long as the core position is well inside the array. The detailed analysis method is described in Nagano *et al.* (1984a). The number of events included for this analysis is 995 457 recorded up to July 1992 for $N_e \geq 10^7$.

The 20km^2 air shower array (hereafter denoted as A20, Teshima *et al.*, 1986) consists of 19 scintillation counters of 2.2m^2 area each and 4 of 1.0m^2 area each, which are distributed with about 1km separation between detectors. A1 is situated at the eastern corner of A20. This array has operated since December 1984. Data are recorded whenever 6 or more neighboring detectors out of 23 are triggered, and at the same time data from all detectors in A1 are recorded. From February 1990, this array has been part of the Akeno Branch of the Akeno Giant Air Shower Array (AGASA). In this analysis, the Akeno Branch data is added to that of A20 up to July 1992 and the number of events analysed is 43 482. Eight muon detectors of A1 are used for showers in A20.

AGASA, which has been operated since February 1990, consists of 111 scintillation counters of 2.2m^2 area each deployed over a surface area of 100km^2 (hereafter denoted as A100, Chiba *et al.* 1992). Fourteen muon detectors have been available for AGASA

1. Introduction

Muons of energies around 1GeV are one of the important observables in extensive air showers (EAS) produced by very energetic cosmic rays. In order to determine the energy and the arrival direction of cosmic rays of the highest energies, the Sydney group (Winn *et al.*, 1986) has measured muons of threshold energies 0.75GeV in large EAS. This is because the total number of muons (N_μ) attenuates with atmospheric depth more slowly than that of electrons (N_e) after their maximum development and hence N_μ may be a good energy estimator of primary cosmic rays. The disadvantage in using muons as an energy estimator is that N_μ is much less than N_e and hence a larger detector is required in measuring muons, and that N_μ is rather sensitive to the model of hadronic interactions and the chemical composition of primary cosmic rays. The latter characteristics are, however, considered to be an advantage to investigate the models of hadronic interactions beyond accelerator energies and the chemical composition in the high energy region where direct measurements of charges of primary particles are not attainable, so long as the primary energy is estimated by other components.

Recently the Yakutsk group claimed that the slope of the lateral distribution of muons changes significantly above 5×10^{18}eV from that at lower energies and that the attenuation length of N_μ becomes very large, while N_e attenuates rather quickly in this energy region. For example, their largest energy event, which has an energy of about 1.2×10^{20}eV and a zenith angle of 60°, consists almost entirely of muons of energies above 2GeV at ground level. The group considers that this may be connected with some new process of EAS development (Glushkov *et al.*, 1994) in the energy region above 5×10^{18}eV. In order to argue the energy and zenith angle dependence of the lateral distribution, careful study is required so as not to bias the distance range concerned, which depends heavily on the primary energy.

The Fly's Eye group reported the energy dependence of the depth at shower maximum (X_{max}) and showed that a change of elongation rate ($\frac{\delta X_{max}}{\delta ln E}$) is observed around 3×10^{17}eV (Bird *et al.*, 1993). They claim that this change is evidence for a change from a predominantly heavy to a predominantly light composition around that energy. Since N_μ from heavy primaries is larger than that from light ones for equal primary energies, the ratio of the muon component to the electron component may become smaller in this energy range than the extrapolated value from the lower energy region, if the Utah group's finding is due to a change in the primary composition.

The observations of extremly high energy cosmic rays well beyond the proposed 2.7K cutoff of the primary spectrum (Bird *et al.*, 1994, Hayashida *et al.*, 1994) has stimulated plans for the next generation experiment. The construction of a giant air shower array of 5000km^2 area is being planned (Cronin, 1992). In this array, a sandwich detector, which consists of a sandwich of a thin top detector, a lead sheet ($\sim$2 r.l.), and a thin bottom detector, is proposed to separate the muonic component from the electromagnetic component far from the core. To design this kind of detector, the determination of the lateral distributions of electrons and muons at core distances exceeding 1000m is necessary.

The characteristics of the muon component are described here using data obtained by three different experiments at Akeno. In Section 2, details of the experiments at the Akeno Air Shower Observatory are described. In Section 3, the lateral distribution of muon densities (ρ_μ) is examined between $10^{16.5}$eV and $10^{19.0}$eV for both vertical and inclined showers with core distances $\leq$800m. The lateral distribution of ρ_μ for core distances exceeding 800m is derived in comparison with that for core distances

Muons ($\geq$1GeV) in Large Extensive Air Showers of Energies between $10^{16.5}$eV and $10^{19.5}$eV Observed at Akeno

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Abstract. The properties of muons ($\geq$1GeV) in giant air showers between $10^{16.5}$eV and $10^{19.5}$eV are measured by the Akeno 1km^2, 20km^2 and 100km^2 air shower arrays. The lateral distribution of muons is well fitted by the formula given by Greisen with parameters $\beta = 2.52 \pm 0.02$ and $\log(R_0) = (0.58 \pm 0.04)(\sec\theta - 1) + (2.39 \pm 0.05)$ between $10^{16.5}$eV and $10^{19.0}$eV within 800m from the air shower axis where θ is the zenith angle of its arrival direction. The relation between total number of muons (N_μ) and total number of electrons (N_e) derived using these parameters is expressed on average as $N_\mu = (2.6 \pm 1.3)\times 10^{5.0+a}(N_e/10^7)^b$ where a=(1.07$\pm$0.13)(secθ–1) and b=(0.77$\pm$0.02)–(0.17$\pm$0.02)(secθ–1). The lateral distribution of muons becomes steeper than that expected from the extrapolation of the above formula for core distances exceeding 800m and is expressed as $\rho_\mu = N_\mu(C'_\mu/R_0^2)\ r^{-0.75}\ (1+r)^{-\beta}\ \{1+(R/800m)^3\}^{-\delta}$ (r=R/R_0, where R is a core distance). No deviation from this formula has been observed for $\sec\theta < 2.0$ up to $10^{19.0}$eV. The relation between the muon density ($\rho_\mu(600)$) and the charged particle density on the ground at 600m from the core (S(600)) is studied between $10^{16.5}$eV and $10^{19.0}$eV. A systematic change in the chemical composition of cosmic rays from a predominantly heavy to a predominantly light composition above $10^{17.5}$eV claimed by the Fly's Eye group has not been detected beyond the present experimental uncertainties. The present experiment suggests much smaller rate of change of composition between $10^{17.5}$eV and $10^{18.5}$eV than that from the Fly's Eye experiment.

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