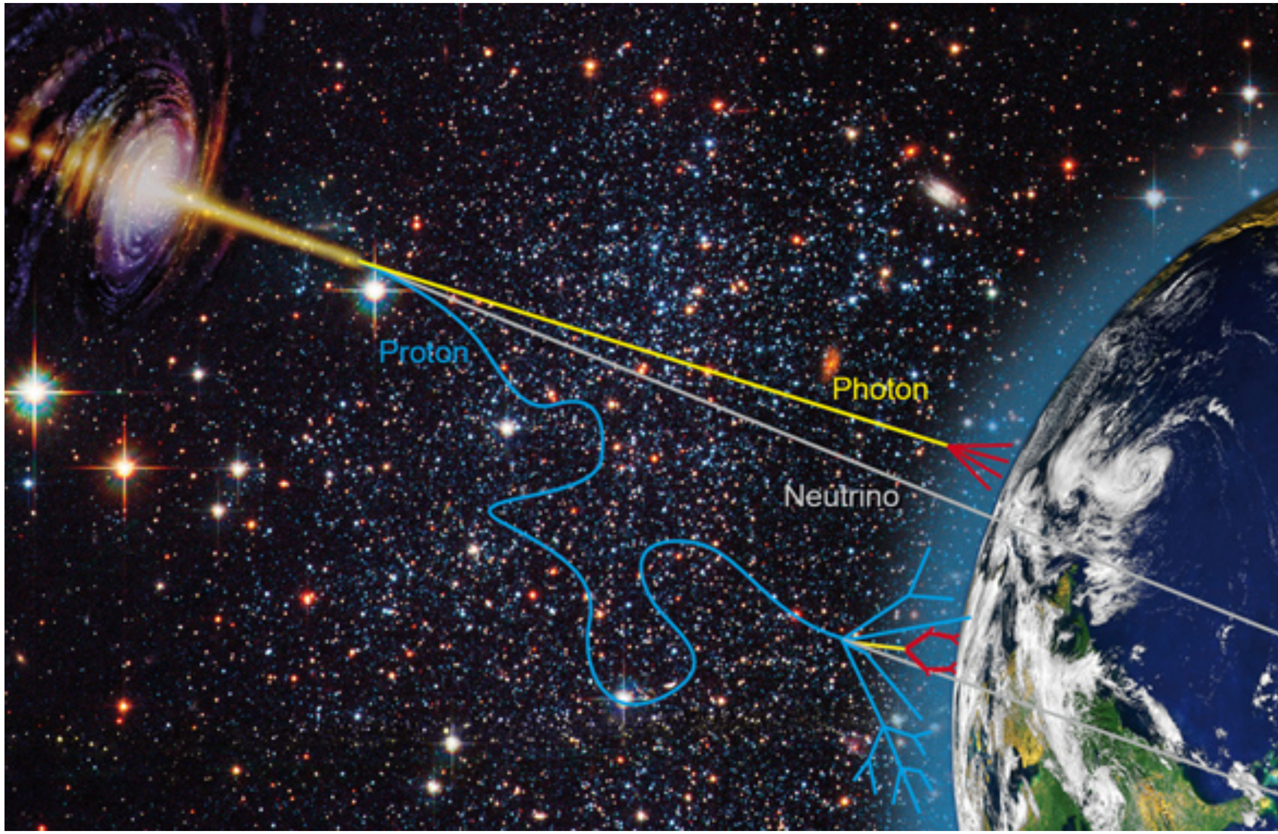
The dawning of electroweak astronomy: interpreting electromagnetic + neutrino emission from blazars

Susumu Inoue (iTHEMS, RIKEN)
on behalf of many collaborators

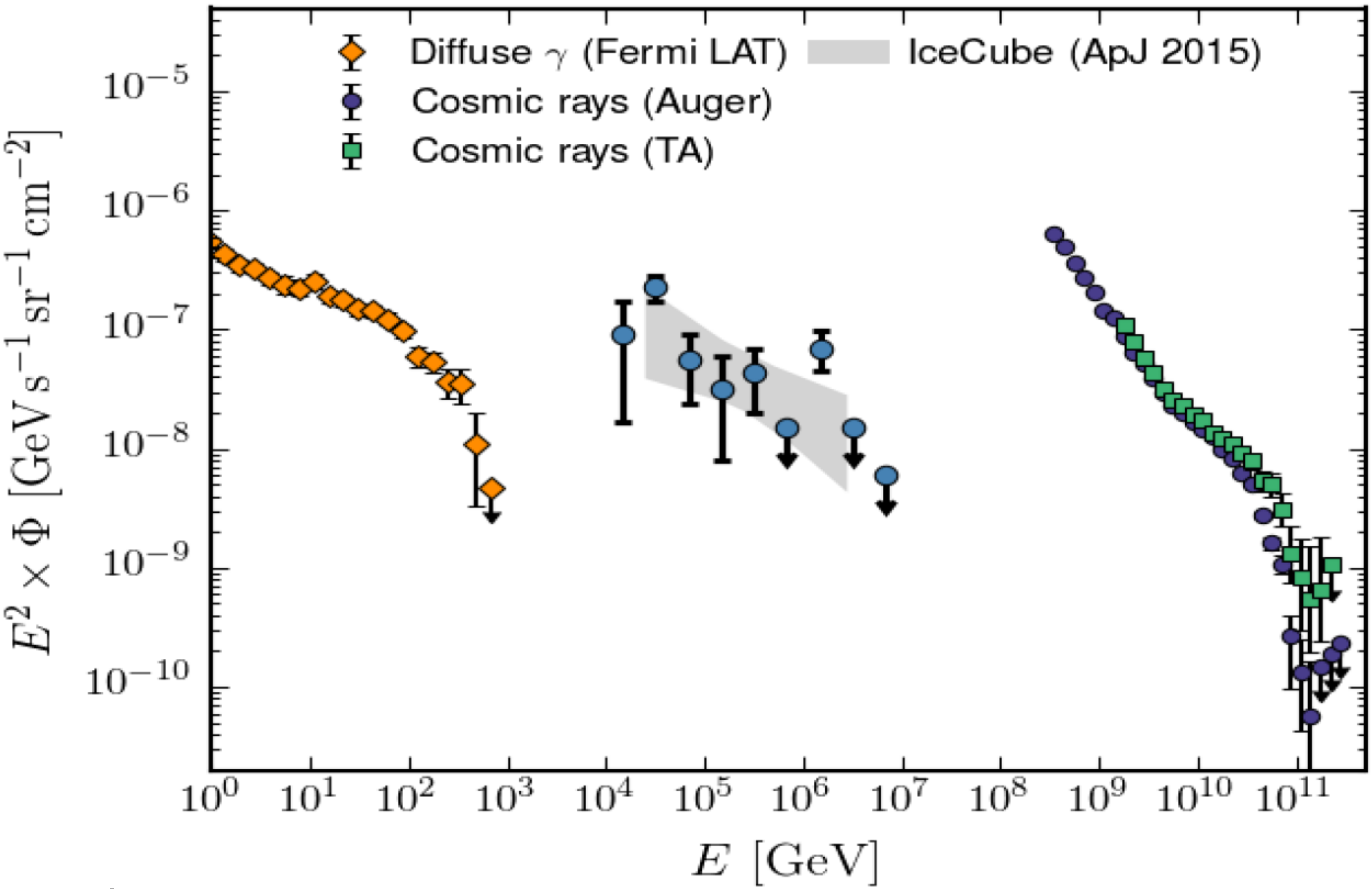


outline

1. introduction
2. electroweak observations of TXS 0506+056 (IC-170922A)
   IceCube, Fermi, MAGIC+, 2018, Science 361, eaat1378
3. interpretation via pγ scenarios (one-zone)
   a. internal photons as targets Cerruti, Zech, … SI+; 1807.04335
   b. "external" photons: jet-sheath (structured jet) scenario
      MAGIC Coll., ApJ 863, L10 (1807.04300)
   c. external photons: radiatively inefficient accretion flow
      Righi, Tavecchio, SI; 1807.10506
   d. brief comparison with other work Keivani, Murase+
4. summary

1. introduction

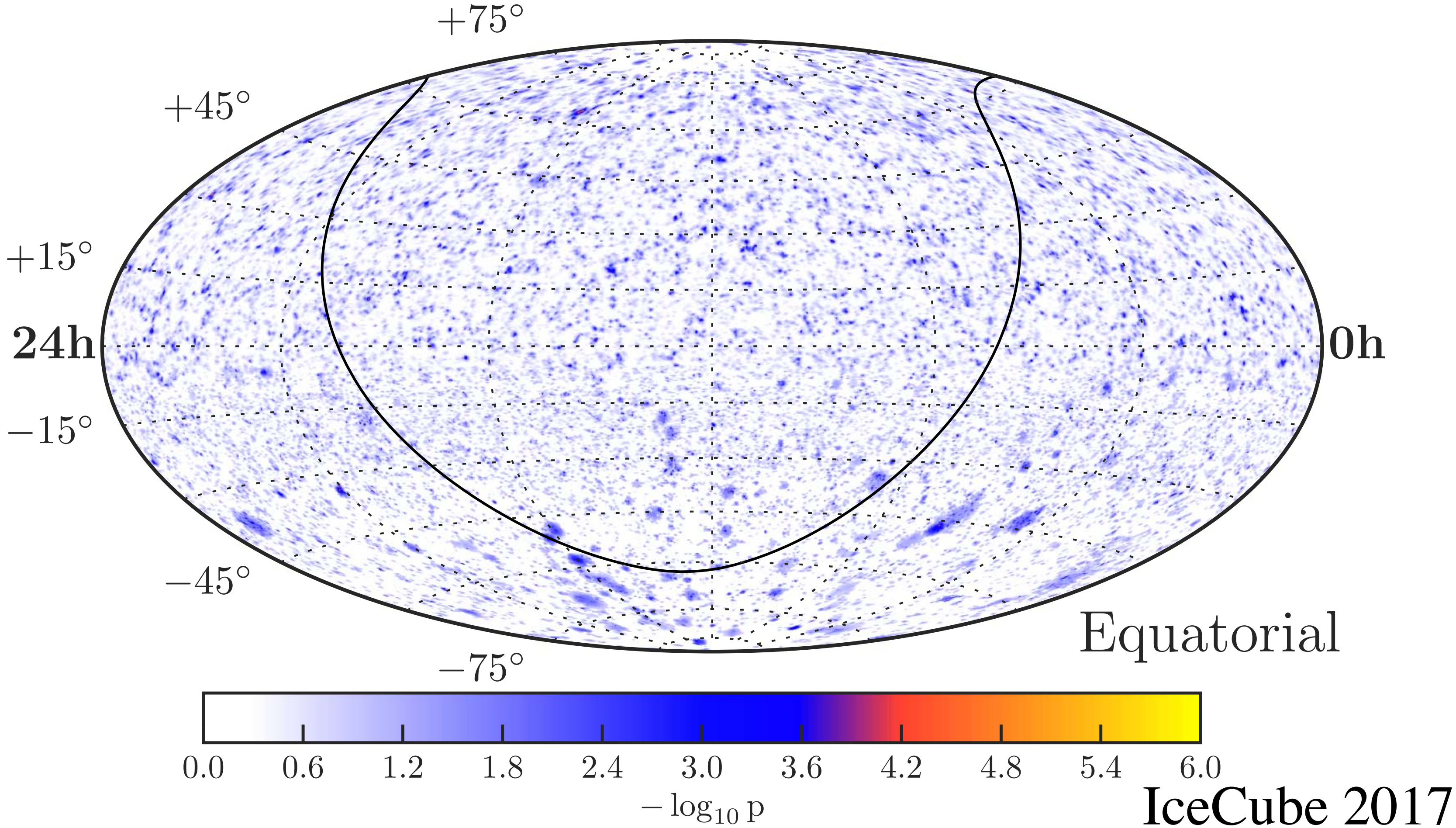
diffuse high-energy neutrinos vs EGB, UHECRs



$p_{CR}+p_{gas}$
$p_{CR}+\gamma_{LE}$ $\rightarrow N+\pi^0, \pi^{+-}$

$\pi^{+-}\rightarrow\mu^{+-}+\nu\rightarrow e^{+-}+3\nu$
$\pi^0\rightarrow 2\gamma$

diffuse high-energy neutrinos: source search

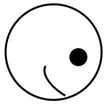


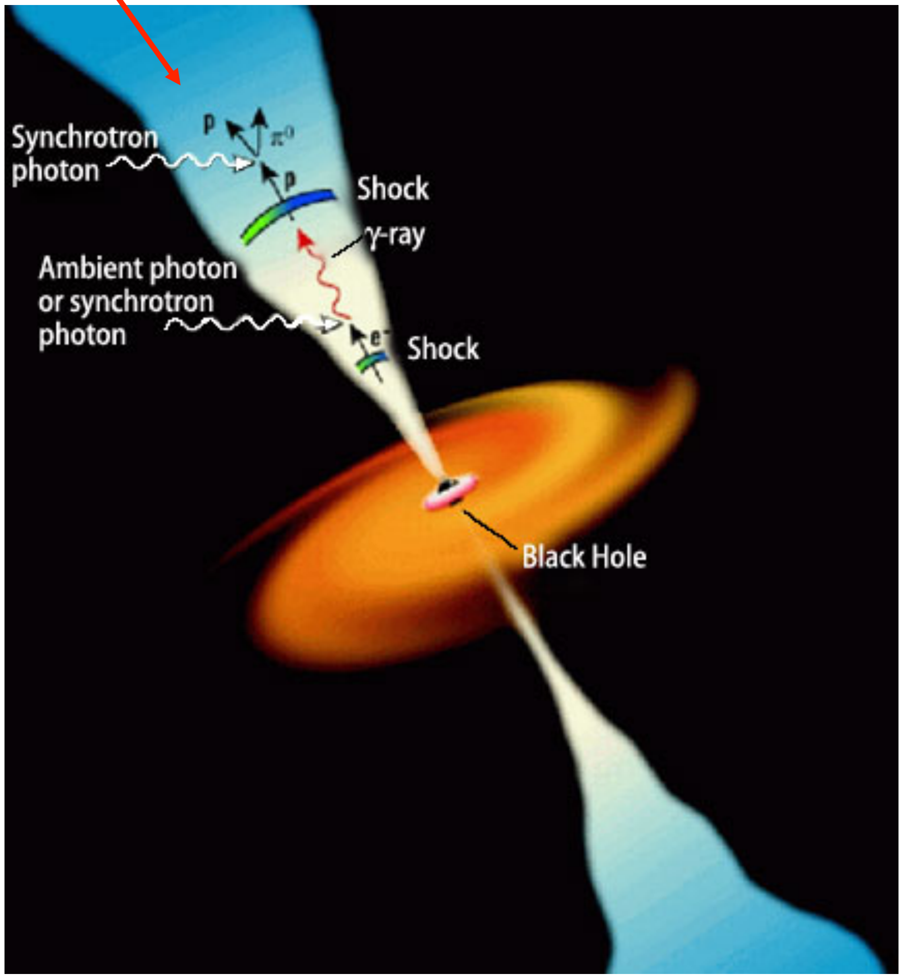
no evidence for source clustering, consistent with isotropy

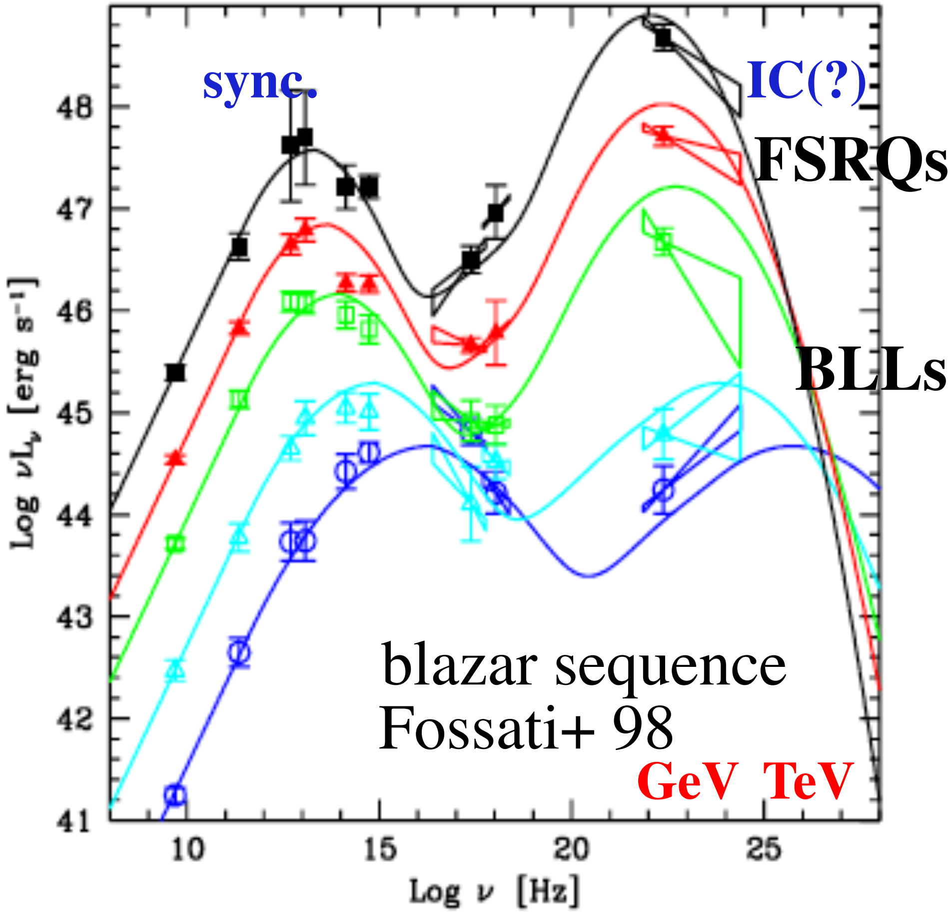
Figure 7: Map of p-values representing the local probability that an excess of events at a given position in the sky is due to a fluctuation of the expected background [47].

blazars

AGN with relativistic jet viewed near-axis
spectra consistent with mainly electron sync.+IC

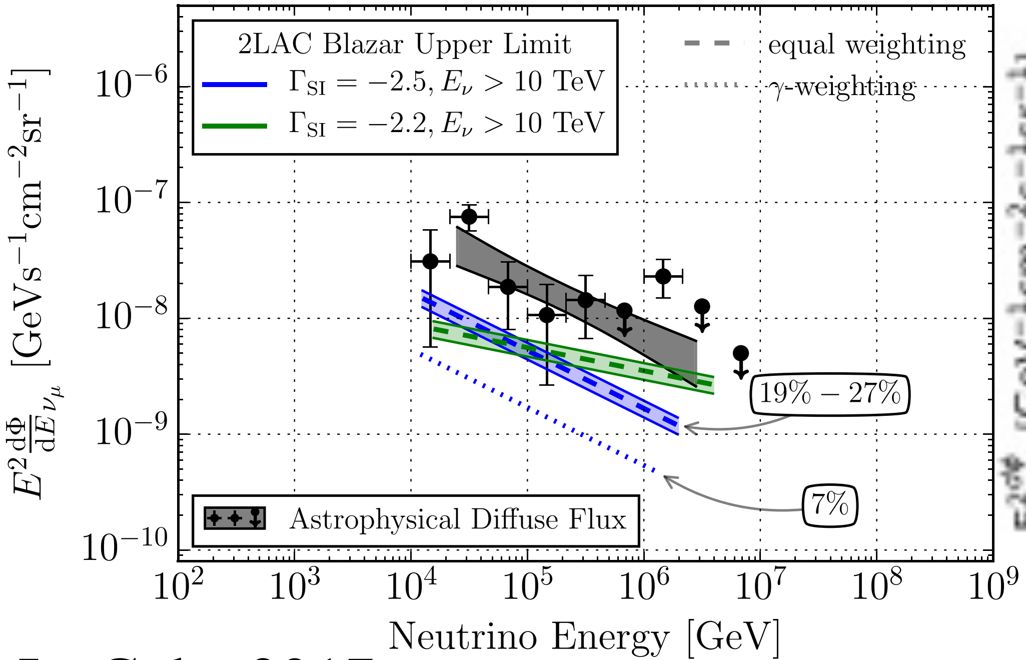




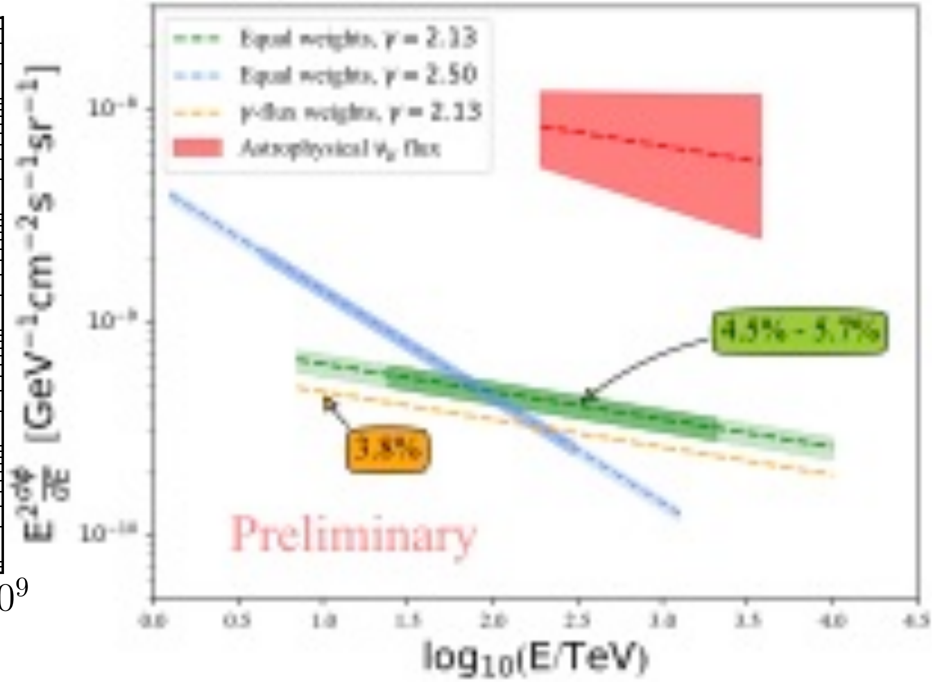


flat spectrum radio quasars (FSRQs): high L, strong emission lines
BL Lac objects (BLLs): low L, weak emission lines

diffuse high-energy neutrinos: constraints on blazar contribution



IceCube 2017



IceCube, ICRC 2017

2LAC: blazar contribution <7-27% of diffuse ν flux
2FHL: blazar contribution <4-6% of diffuse ν flux

blazars strongly constrained as main sources of diffuse HE ν

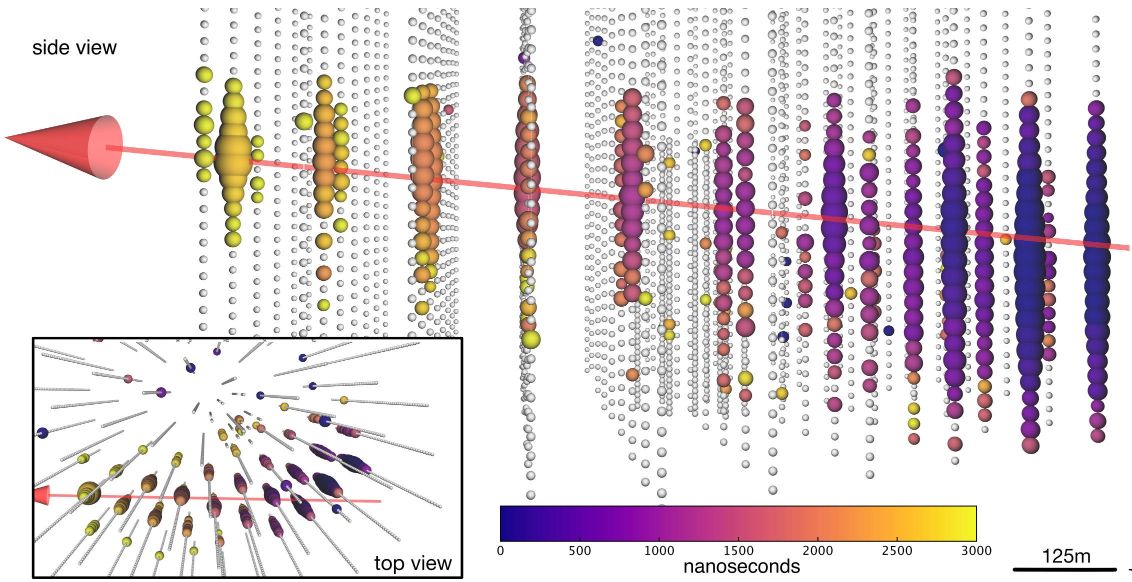
see however Palladino+ 1806.04769

even if so, some individual blazars may still be detectable

2. electroweak observations of TXS 0506+056 / IC-170922A

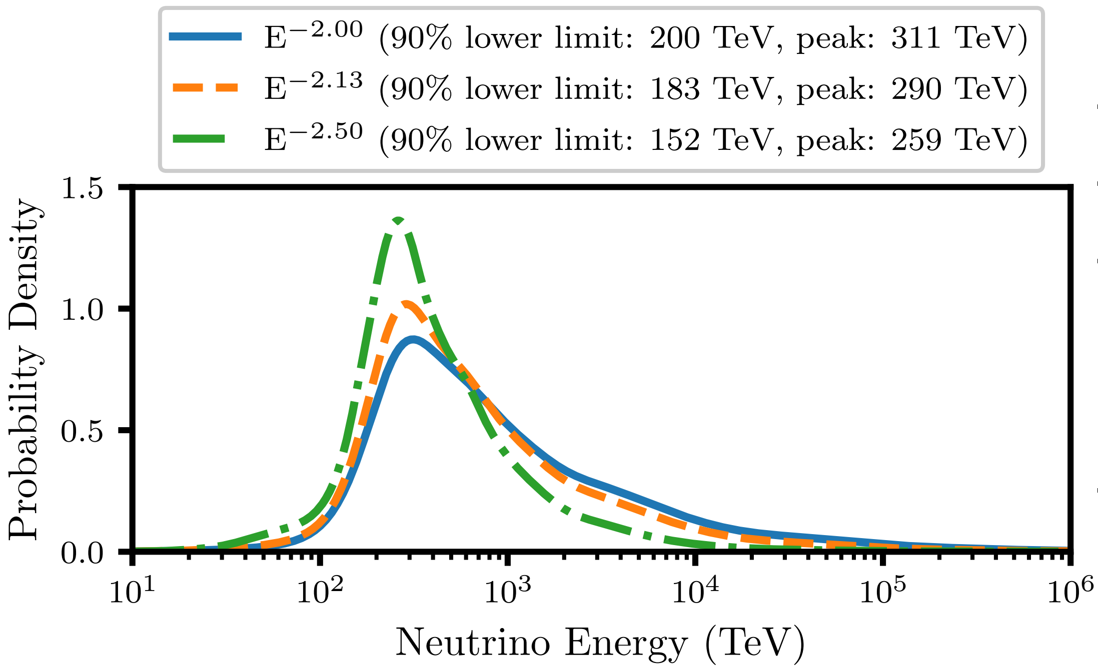
ν observation of IC-170922A

IceCube, Fermi, MAGIC+, 2018, Science 361, eaat1378

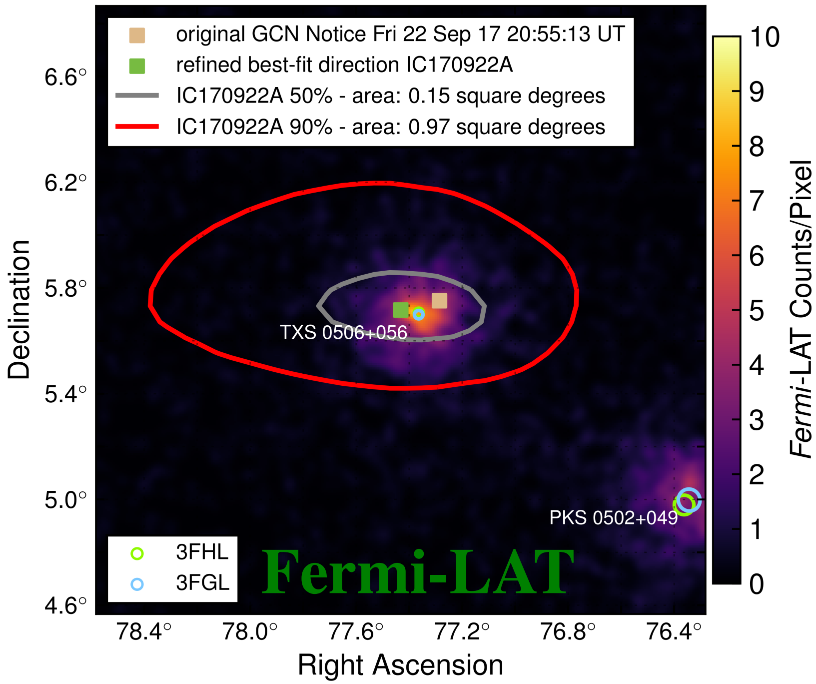


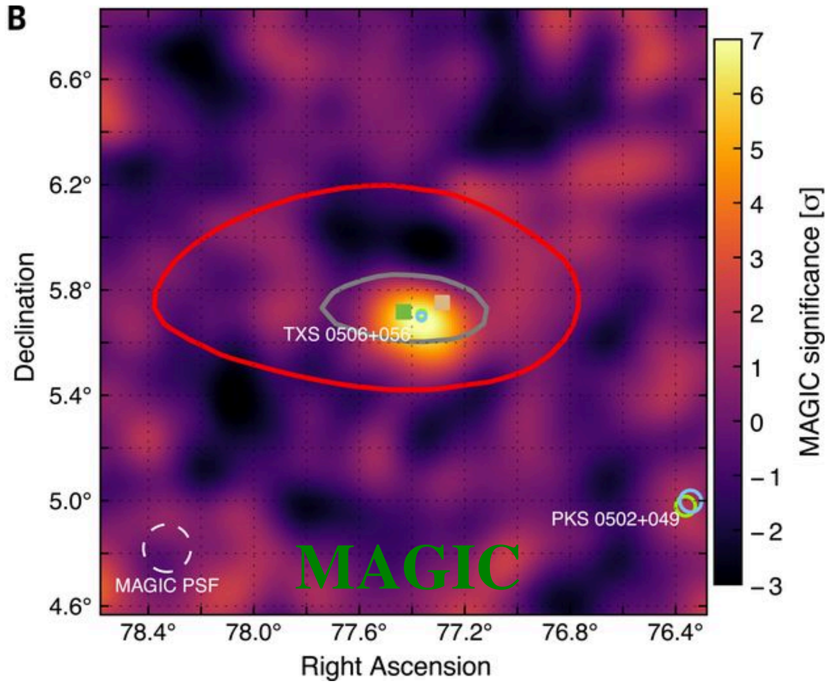
- EHE event: 56.5% probability of being astrophysical ν
- alert after 43s
- well localized, ~<1 deg
- E_ν~290 TeV (183 TeV - 4.3 PeV 90% C.L.) assuming -2.13 spectrum

-> possible cosmic proton accelerator with E_p>~$20E_\nu$~several PeV



ν + EM observations of IC-170922A / TXS 0506+056





Fermi-LAT:

- coincident with blazar TXS 0506+056 in bright state (0.5 yr-long)
- significance of association ~3σ

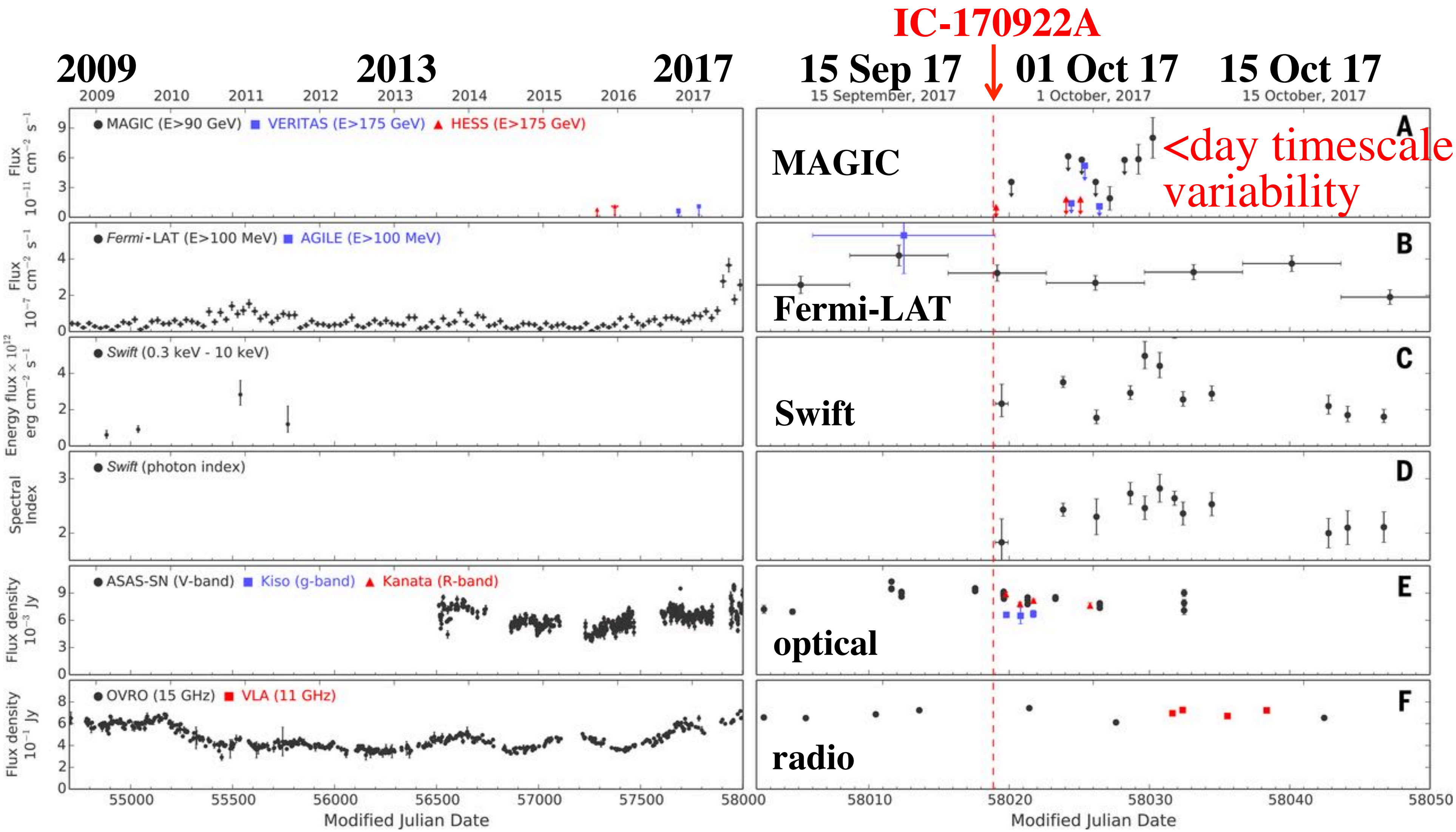
-> possible source of possible astrophysical high-energy neutrino

MAGIC:

- ~6σ detection, <day timescale flaring
- steep spectrum (Γ~-3.5 - -4.0) up to 400 GeV

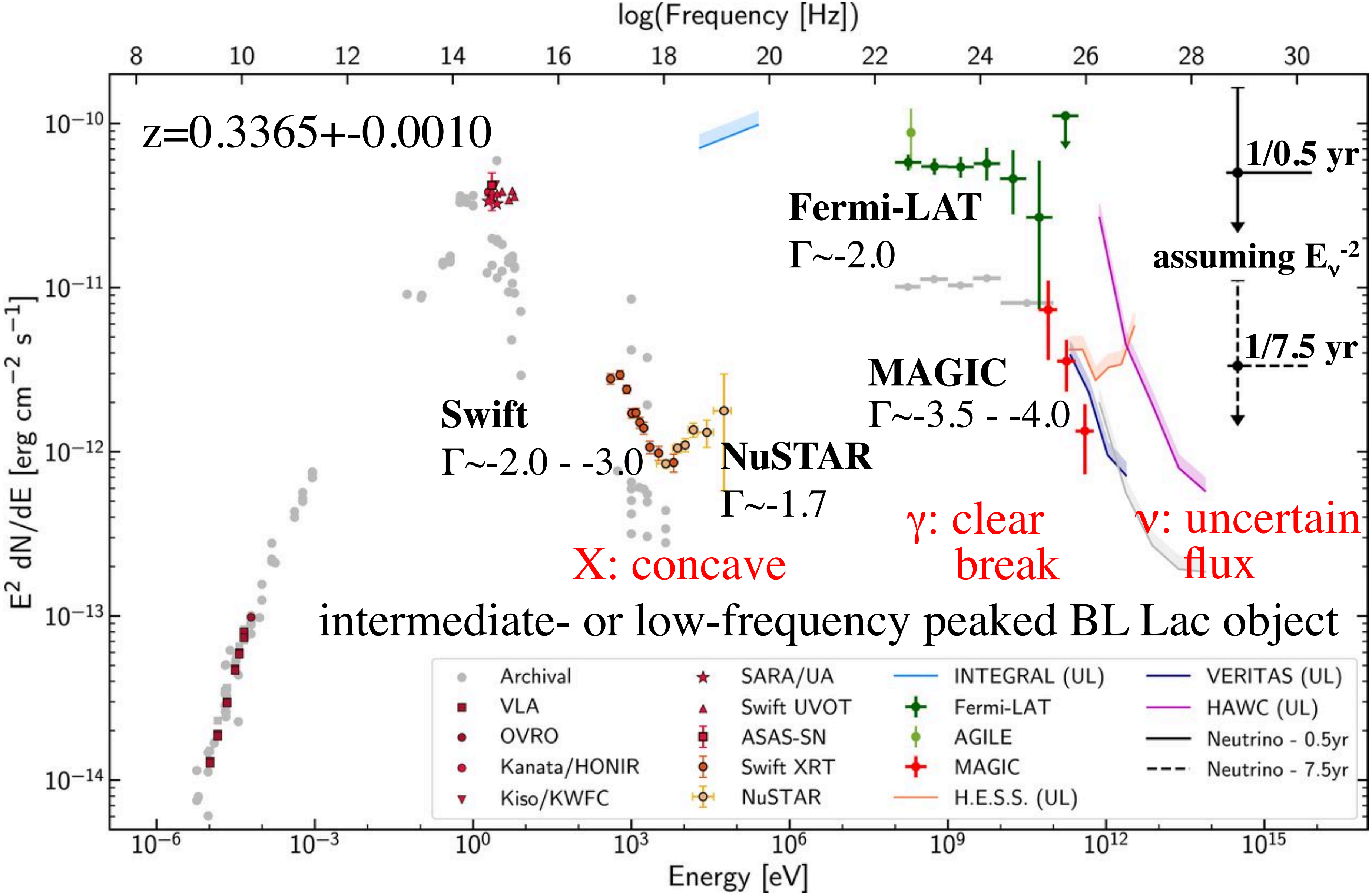
-> crucial contraints on physical conditions of source

ν + EM observations of IC-170922A / TXS 0506+056



IceCube, Fermi, MAGIC+, 2018, Science 361, eaat1378

ν + EM observations of IC-170922A / TXS 0506+056



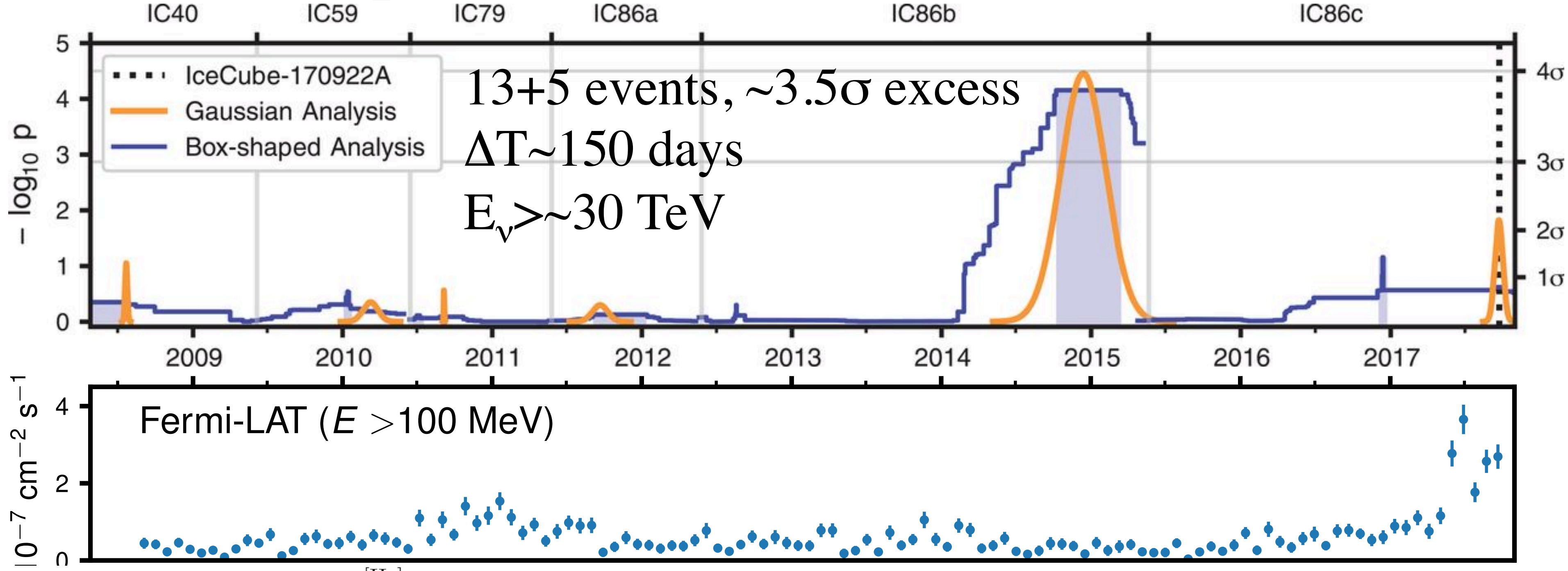
IceCube, Fermi, MAGIC+, 2018, Science 361, eaat1378

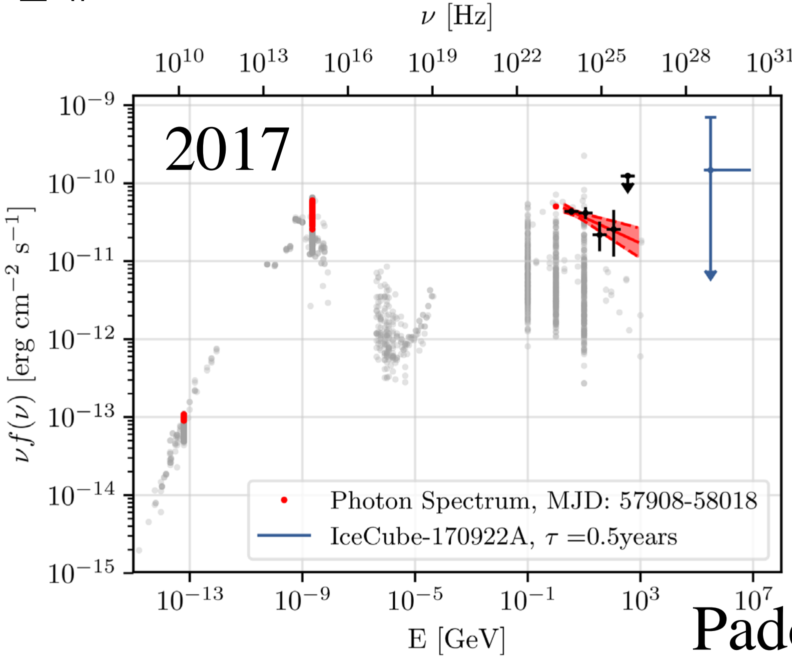
Redman’s theorem

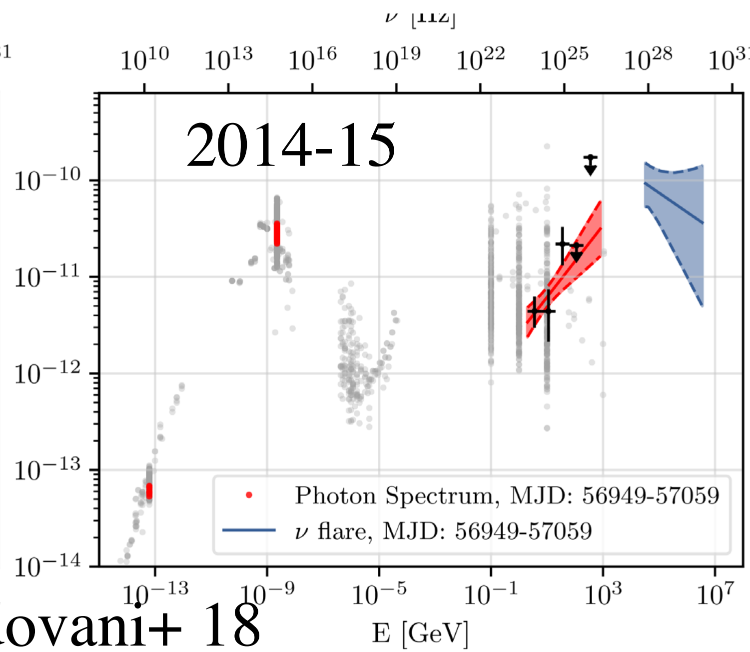
“A competent theoretician can fit any given theory to any given set of facts.”

2014-2015 (orphan) neutrino flare

IceCube, 2018, Sci. 361, 147







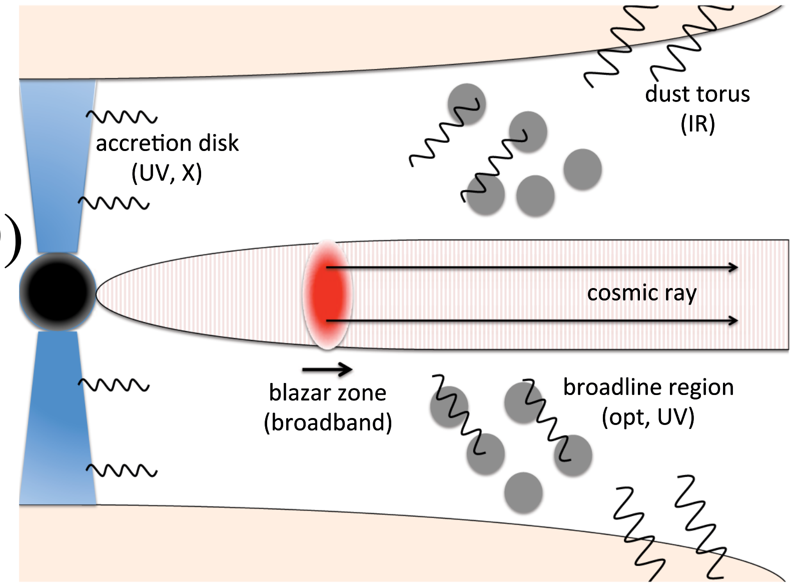
Padovani+ 18

viable interpretation via pp? modeling ongoing, stay tuned

3. interpretation via pγ scenarios

neutrino emission from blazars

- $p\gamma$ generally favored over pp in AGN jets
- target γ $\varepsilon'_{\gamma} >\sim 20 m_{\pi} m_p c^4 / E_{\nu} \delta^{-1}$
  ~ 0.4 keV $(E_{\nu}/300 \text{ TeV})^{-1}$ $(\delta/20)$

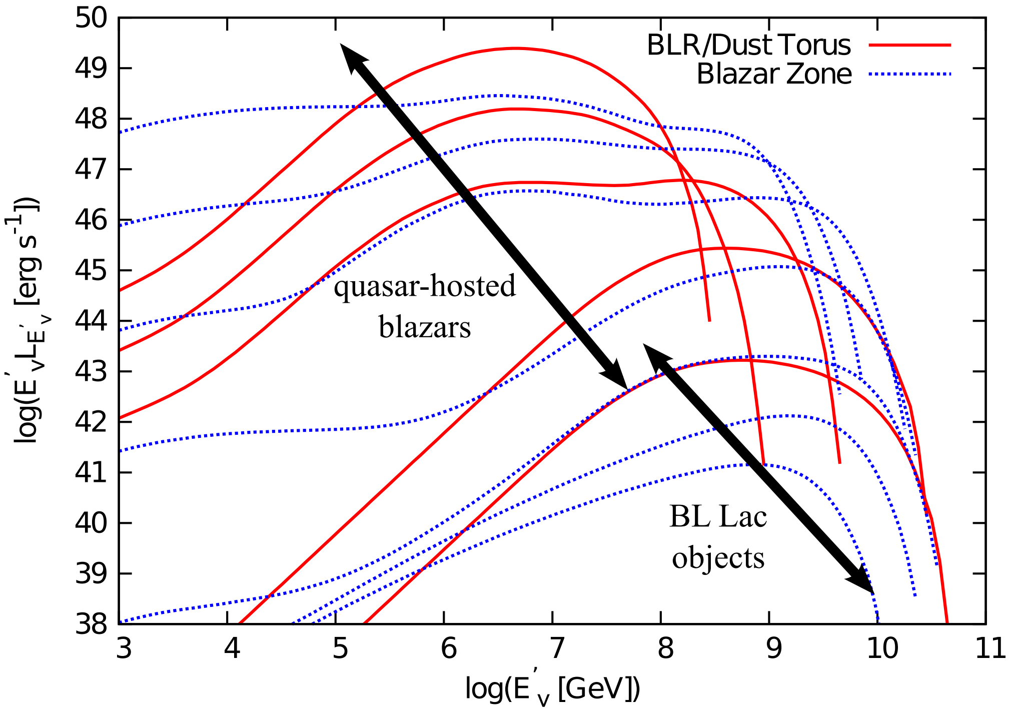


Murase+ 14

- unlike FSRQs, BL Lacs thought to:
  lack bright external γ fields,
  have low internal sync. γ fields
  -> PeV ν production inefficient?

- one-zone models (e+p co-accel.)
  with internal syn. targets only:
  very high L_p required
  Cerruti, Zech, Boisson, Emery, SI, Lenain, 1807.04335

  see also Gao+ 1807.04275

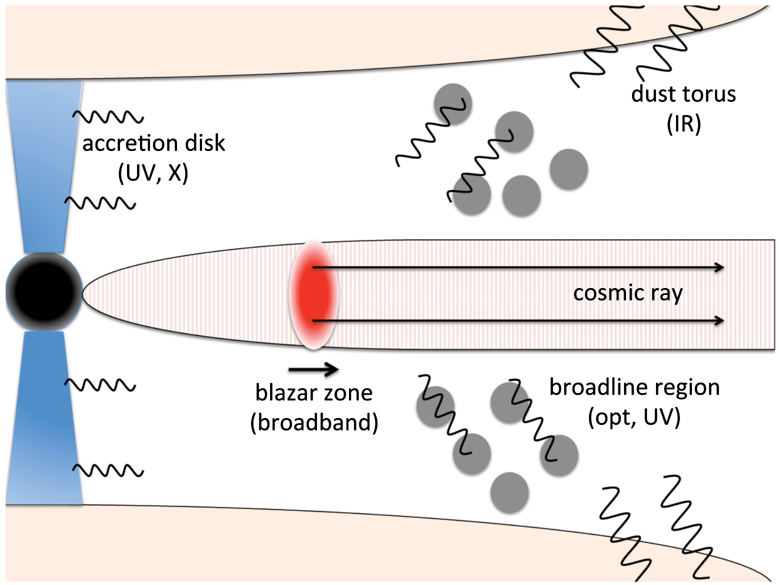


neutrino emission from blazars

- enhanced pγ efficiency via external γ fields in BL Lacs?

I. sync. from sheath in structured jets MAGIC Coll. 1807.04300

II. radiatively inefficient accretion flows (RIAFs) Righi+ 1807.10506

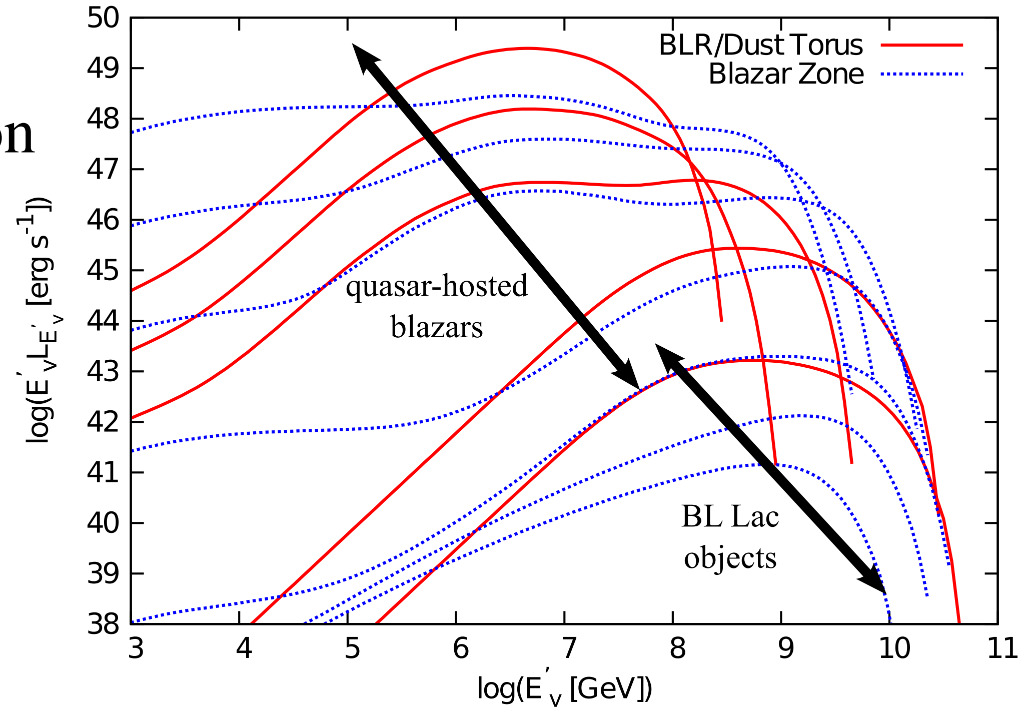


Murase+ 14

focus on one-zone models (electron+proton co-accelerated)

- questions

1. accompanying hadronic emission consistent with observed SED?
2. accompanying γγ absorption consistent with observed SED?
3. role of external Compton relative to SSC?
4. proton maximum energy (UHECR accelerator)?



3a. pγ scenarios with internal photons only model description

Cerruti, Zech, …
SI+ 1807.04335

leptonic part

- emission region: spherical with radius R, magnetic field B, Doppler factor δ
- electron distribution: broken power-law $\gamma_{e,min}$, $\gamma_{e,max}$, α_{e1},α_{e2}
- leptonic emission: synchrotron, SSC

hadronic part   follow Cerruti+ 15, Zech+ 17

- proton distribution: power-law $\gamma_{p,min}=1$, $\gamma_{p,max}$ (or η), $\alpha_{p1}=\alpha_{e1}$
- hadronic emission

$p+\gamma_{LE}\rightarrow N+\pi^0, \pi^{+-}$ photo-meson   SOPHIA: Mücke+ 02,03

$\pi^{+-}\rightarrow\mu^{+-}+\nu\rightarrow e^{+-}+3\nu$ $\pi^0\rightarrow 2\gamma$

$\mu^{+-}+B\rightarrow\mu^{+-}+\gamma$ muon synchrotron

$p+\gamma_{LE}\rightarrow p+e^+e^-$ photo-pair (Bethe-Heitler)   Kelner & Aharonian 08

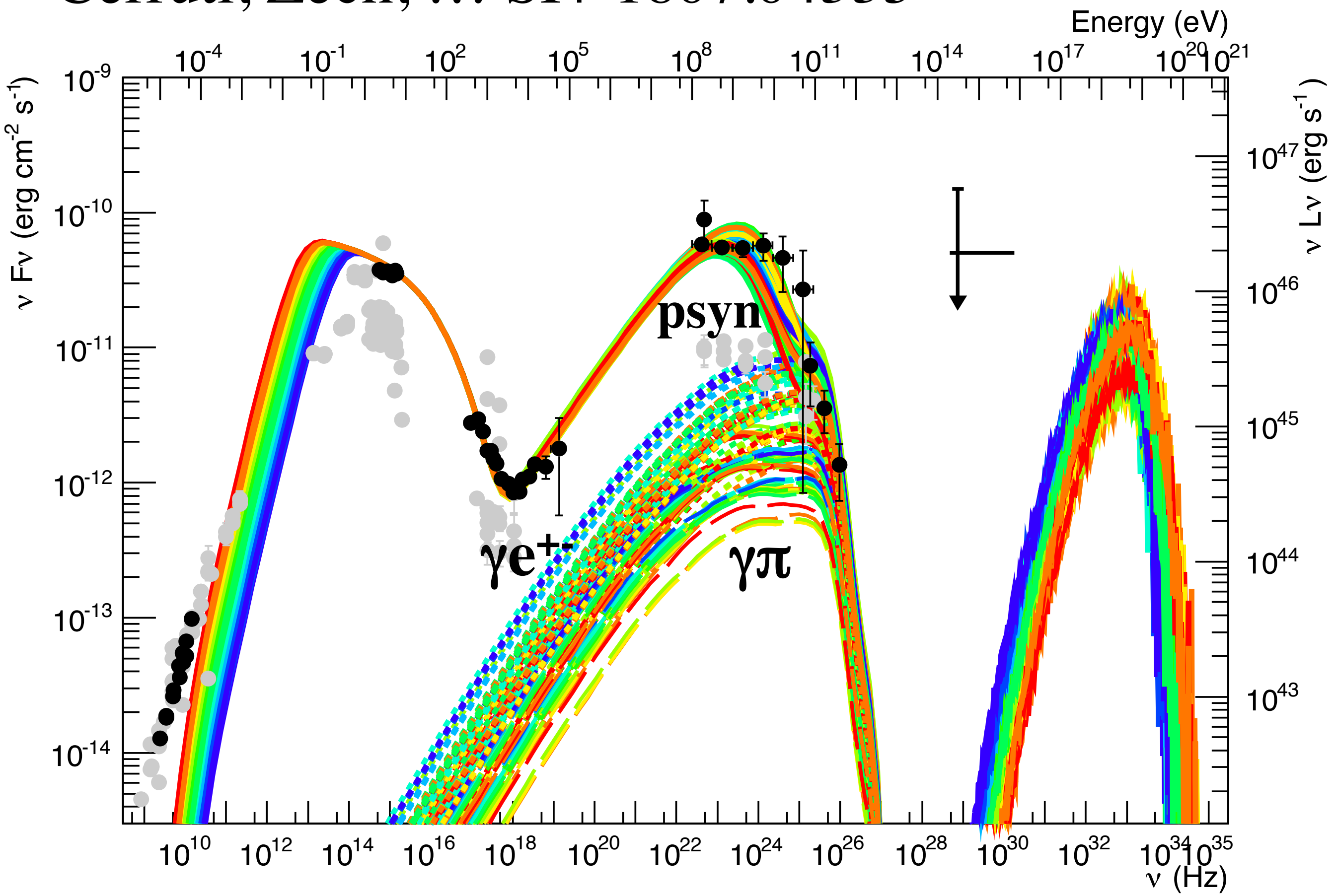
$\gamma+\gamma_{LE}\rightarrow e^+e^-$

$e^+e^-+B\rightarrow e^+e^-+\gamma$ electron-positron sync. cascade

$p+B\rightarrow p+\gamma$ proton synchrotron

proton sync.-dominant case (+pγ cascade)

Cerruti, Zech, … SI+ 1807.04335

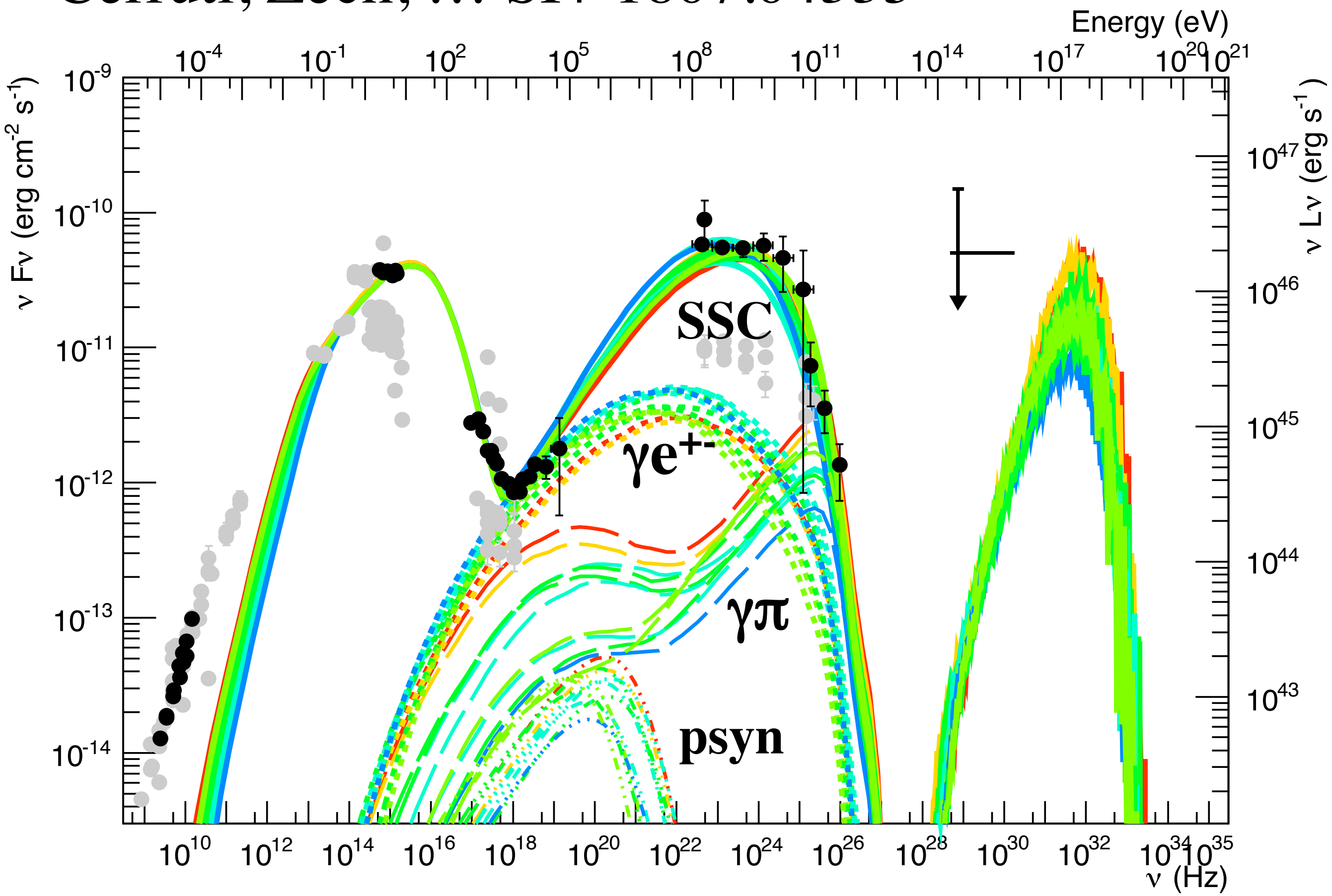


	Proton-synchrotron
z	0.337
δ	35 – 50
R [10^{16} cm]	0.1 – 9.7
$^\star\tau_{\rm obs}$ [days]	0.01 – 1.0
B	0.8 – 32
$^\star u_B$ [erg cm^{-3}]	0.02 – 0.16
$\gamma_{e,\rm min}$	500
$\gamma_{e,\rm break}$	$= \gamma_{e,\rm min}$
$\gamma_{e,\rm max}$ [10^4]	0.6 – 1.0
$\alpha_{e,1} = \alpha_{p,1}$	2.0
$\alpha_{e,2} = \alpha_{p,2}$	3.0
K_e [cm^{-3}]	$6.3 - 9.1 \times 10^3$
$^\star u_e$ [10^{-5} erg cm^{-3}]	0.4 – 15.1
$\gamma_{p,\rm min}$	1
$\gamma_{p,\rm break}[10^9]$	$= \gamma_{p,\rm max}$
$\gamma_{p,\rm max}[10^9]$	0.4 – 2.5
η	20 – 50
K_p [cm^{-3}]	$10.4 - 2.0 \times 10^4$
$^\star u_p$ [erg cm^{-3}]	0.7 – 45
$^\star u_p/u_B$	1.0 – 89
$^\star L$ [10^{46} erg s^{-1}]	0.8 – 170
$^\star\nu$ [year^{-1}]	$5.7 \times 10^{-3} - 0.2$
$^\star\nu_{183-4300\,\rm TeV}$ [year^{-1}]	$2.4 \times 10^{-5} - 1.7 \times 10^{-3}$

- EM SED reproduceable with proton synchrotron dominating γ rays, pγ cascade non-negligible at VHE
- BUT neutrino flux too low to be viable
  -> detection of single ν provides crucial discriminant

SSC-dominant case (+pγ cascade)

Cerruti, Zech, … SI+ 1807.04335



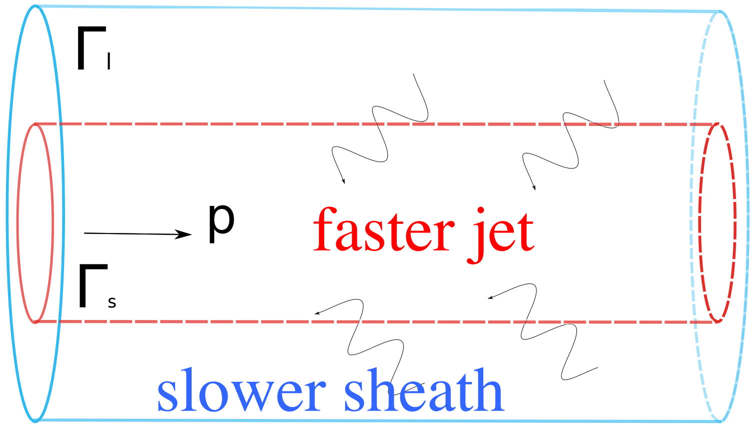
	Lepto-hadronic
z	0.337
δ	30 – 50
R [10^{16} cm]	0.2 – 1.5
$^{\star}\tau_{\rm obs}$ [days]	0.02 – 0.3
B	0.13 – 0.65
$^{\star}u_B$ [erg cm^{-3}]	6.5×10^{-4} – 0.017
$\gamma_{e,\rm min}$	500
$\gamma_{e,\rm break}$	$= \gamma_{e,\rm max}$
$\gamma_{e,\rm max}$ [10^4]	0.8 – 1.7
$\alpha_{e,1} = \alpha_{p,1}$	2.0
$\alpha_{e,2} = \alpha_{p,2}$	3.0
K_e [cm^{-3}]	9.5×10^3 – 2.6×10^5
$^{\star}u_e$ [10^{-5} erg cm^{-3}]	2.2×10^3 – 43×10^3
$\gamma_{p,\rm min}$	1
$\gamma_{p,\rm break}$[10^9]	$= \gamma_{p,\rm max}$
$\gamma_{p,\rm max}$[10^9]	0.06 – 0.2
η	10
K_p [cm^{-3}]	3.5×10^3 – 6.6×10^4
$^{\star}u_p$ [erg cm^{-3}]	100 – 1400
$^{\star}u_p/u_B$	3.9×10^4 – 79×10^4
$^{\star}L$ [10^{46} erg s^{-1}]	35 – 350
$^{\star}\nu$ [year^{-1}]	0.11 – 3.0
$^{\star}\nu_{183-4300\,\rm TeV}$ [year^{-1}]	0.008 – 0.11

- EM SED reproduceable with SSC dominating γ rays,
  non-neglible: pγ cascade at VHE, BH at X-ray
- BUT requires rather extreme parameters, e.g. L_p~10^{48} erg/s, $\gamma_{e,min}$~500
- hard ν spectra also constrained by IC point source search
  -> scenarios with external photons likely more favorable

3b. pγ scenarios with “external” photons from jet sheath

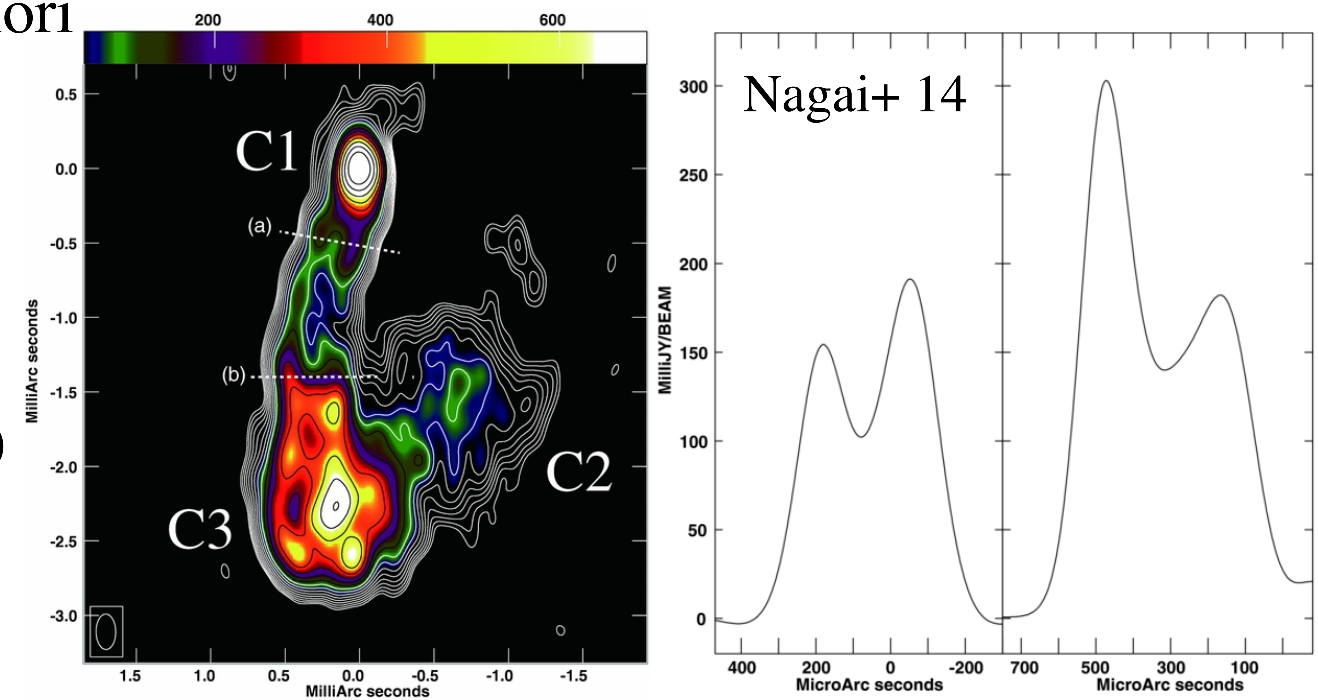
MAGIC Coll. 1807.04300

- jet structure with slower sheath (layer) surrounding faster jet (spine)
  -> supported by observations, numerical simulations
- synchrotron photons from sheath seen Doppler boosted in jet frame
  -> enhanced target for pγ ν production, EC
- caveat: spectrum of sheath radiation not well defined a-priori



c.f. Tavecchio+ 14, 15
Righi & Tavecchio 17

limb-brightened structure in radio galaxies e.g. 3C84 (NGC 1275)



jet-sheath model description

MAGIC Coll. 1807.04300

leptonic part   follow Tavecchio+ 14, 15

- emission region: cylindrical with radius R, length dR=R, magnetic field B, Lorentz factor Γ_j, viewing angle θ_v
- electron distribution: broken power-law $E_{e,min}$, $E_{e,br}$, $E_{e,max}$, s_1, s_2
- photons from sheath with Lorentz factor Γ_s, broken power-law spectrum
- leptonic emission: synchrotron, SSC, EC

hadronic part   follow Böttcher+ 13, Cerruti+ 15

- proton distribution: power-law E_p^{-2} with exp. cutoff E_{pmax}
- hadronic emission

$p+\gamma_{LE} \rightarrow N+\pi^0, \pi^{+-}$ photo-meson

$\pi^{+-} \rightarrow \mu^{+-}+\nu \rightarrow e^{+-}+3\nu$ $\pi^0 \rightarrow 2\gamma$

$\mu^{+-}+B \rightarrow \mu^{+-}+\gamma$ muon synchrotron

Kelner & Aharonian 08

$p+\gamma_{LE} \rightarrow p+e^+e^-$ photo-pair (Bethe-Heitler)

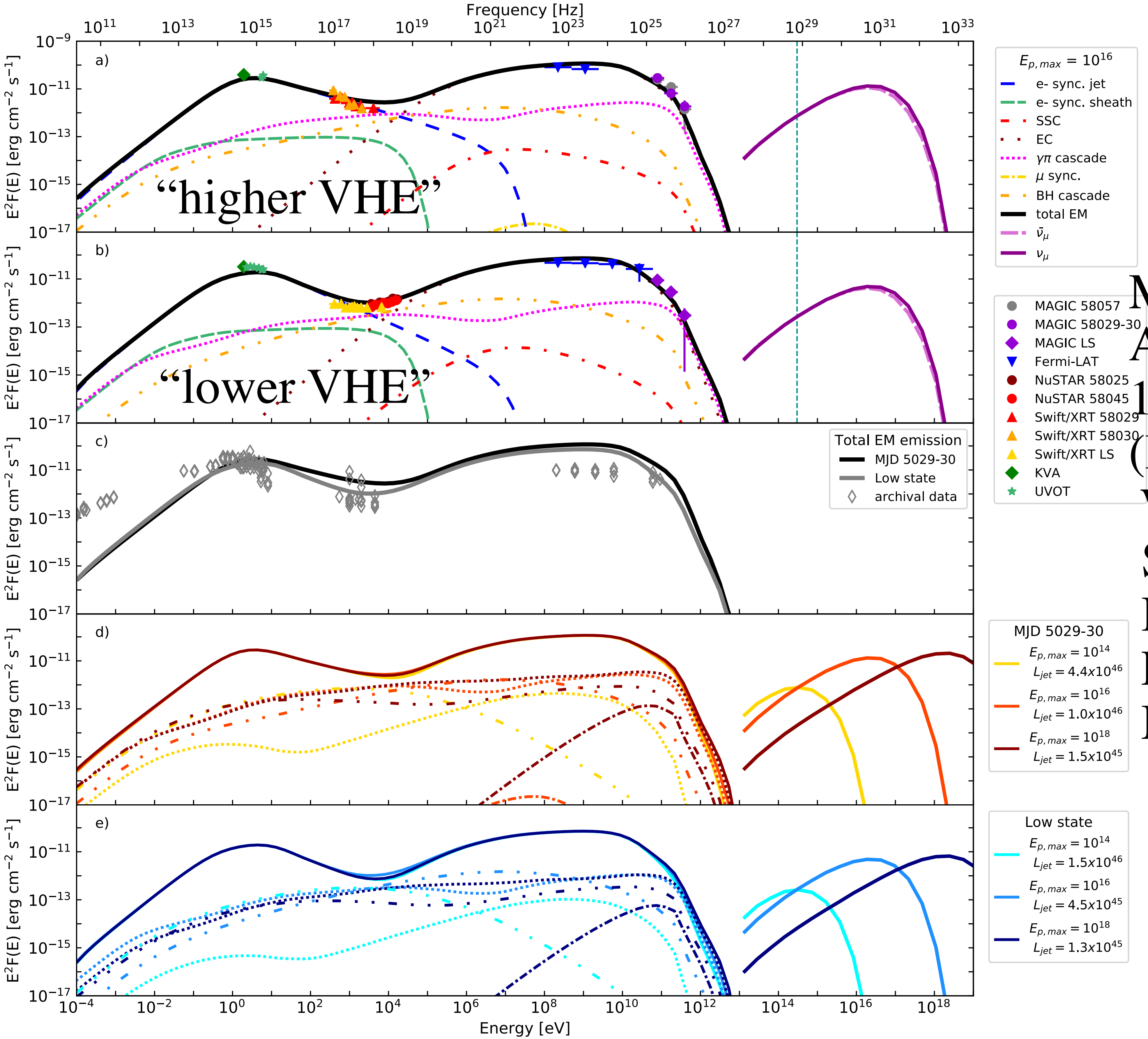
$\gamma+\gamma_{LE} \rightarrow e^+e^-$

$e^+e^-+B \rightarrow e^+e^-+\gamma$

electron-positron sync. cascade

$p+B \rightarrow p+\gamma$ proton synchrotron

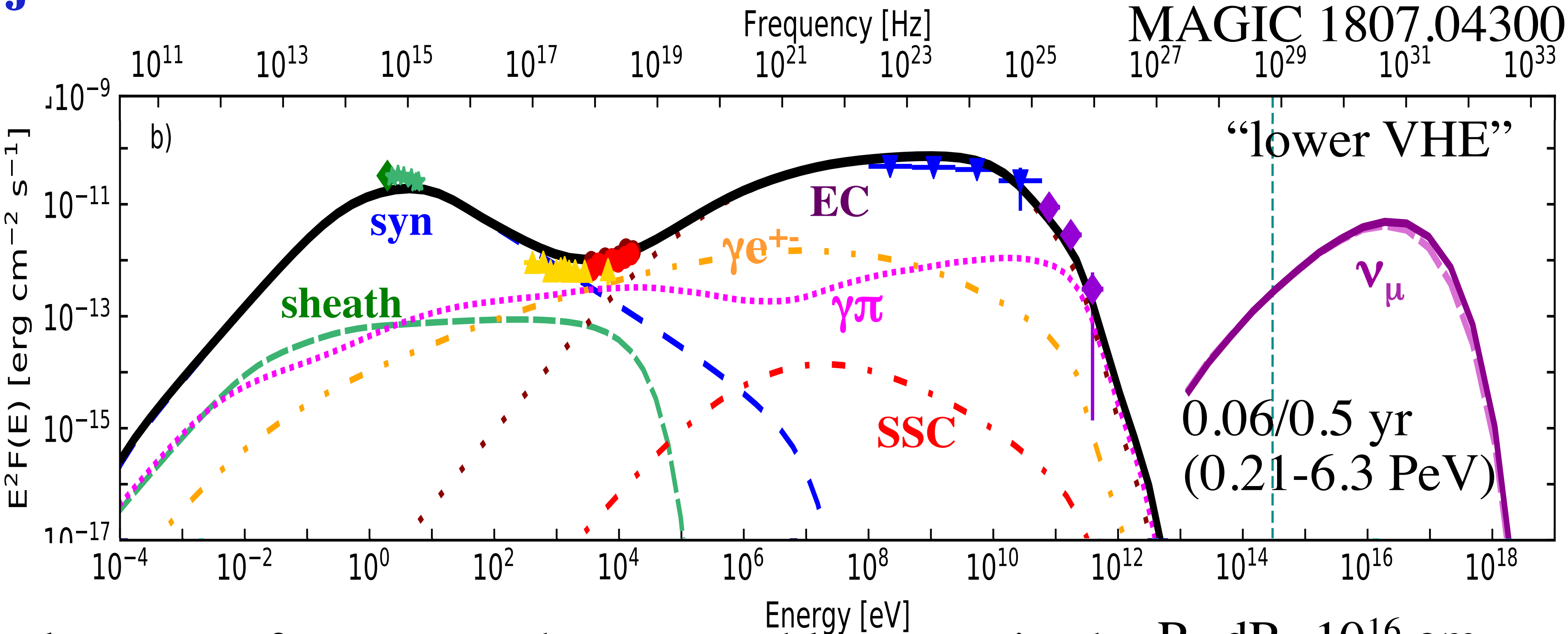
jet-sheath model for electroweak emission



MAGIC Coll.
ApJ 863, L10
1807.04300
(E. Bernardini
W. Bhattacharya
SI
K. Satalecka
F. Tavecchio)
M. Cerruti

jet-sheath model for electroweak emission

MAGIC 1807.04300



- large no. of parameters but reasonably constrained
- SED predominantly leptonic, γ-rays EC (not SSC)
- photopion+photopair cascade subdominant but non-negligible in X (+VHE)
  -> crucial constraint on proton population
- photopion efficiency $f_{p\gamma}(E_p\sim 6\ \mathrm{PeV})\sim 10^{-4}$
  -> $\tau_{\gamma\gamma}(E_\gamma\sim 12\ \mathrm{GeV})\sim 0.1$ -> $\tau_{\gamma\gamma}(E_\gamma\sim 100\ \mathrm{GeV})\sim 1$
  consistent with observed GeV-TeV break

$R=dR=10^{16}$ cm
$B=2.6$ G
$\Gamma_j=22$, $\Gamma_s=2.2$
$\theta_v=0.8°(\delta_j=40)$
$E'_{pmax}=10^{16}$ eV
$L_e=1.6\times10^{42}$ erg/s
$L_p=3\times10^{45}$ erg/s(?)
($L_B=1.2\times10^{45}$ erg/s)

summary of constraints for external photon scenarios

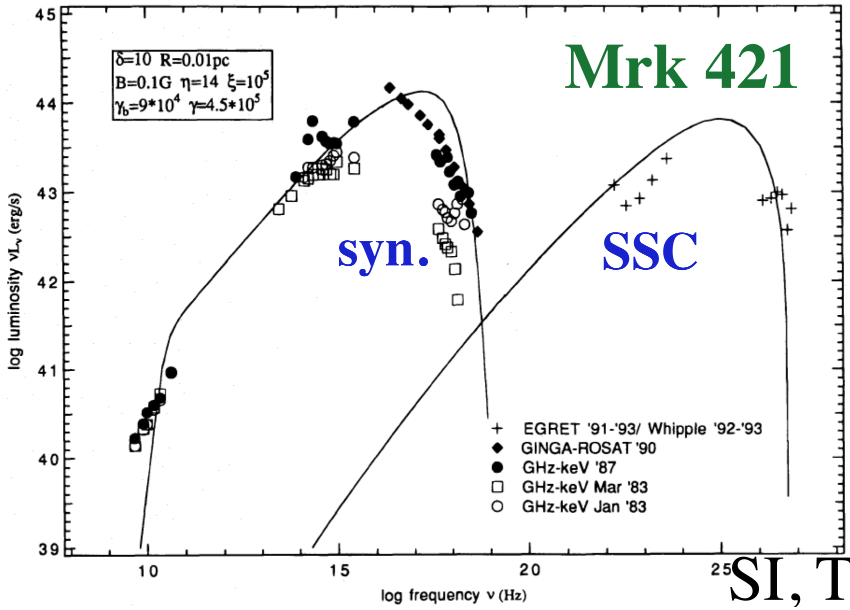
- proton power × soft photon density:
  high enough to explain observed (but uncertain) neutrino flux
- proton power × soft photon density:
  low enough for cascade not to violate X-ray constraints
- soft photon density:
  low enough for $\gamma\gamma$ absorption not to violate VHE constraints
- soft photon spectrum:
  constrained to reproduce γ ray spectrum via external Compton
- proton power:
  low enough to comply with (uncertain) Eddington constraints

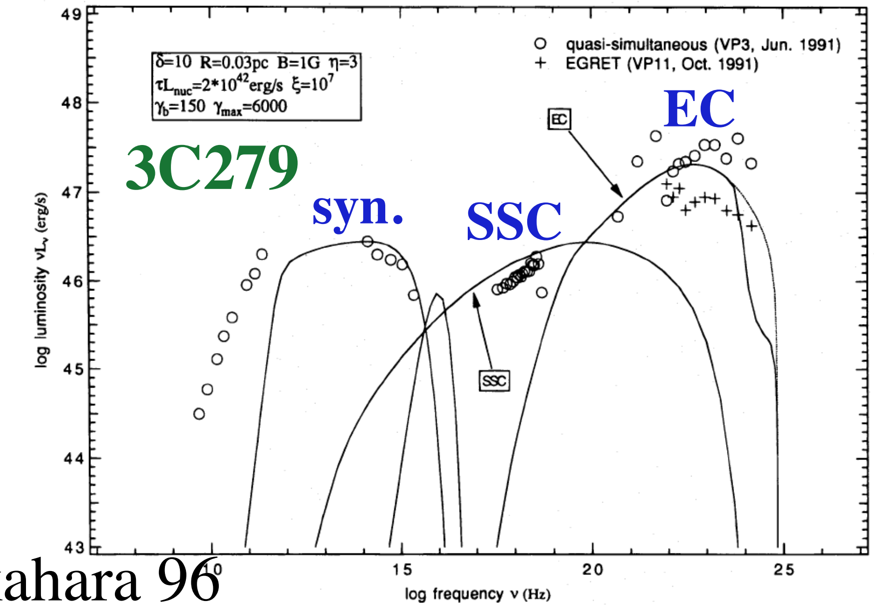
implications of jet-sheath model: jet energy balance

	lower VHE	higher VHE
L_e [erg/s]	1.6×10^{42}	2×10^{42}
L_p [erg/s]	3×10^{45}(?)	8×10^{45}(?)
L_B [erg/s]	1.2×10^{45}	1.2×10^{45}
U_p/U_e	1700	3600

- SSC subdominant -> $U_B \sim (<) U_p$ near equipartion, U_e subdominant
  c.f. Ghisellini+ 05 potentially consistent with B-dominant jets
- proton/electron $U_p/U_e (>) \sim m_p/m_e$

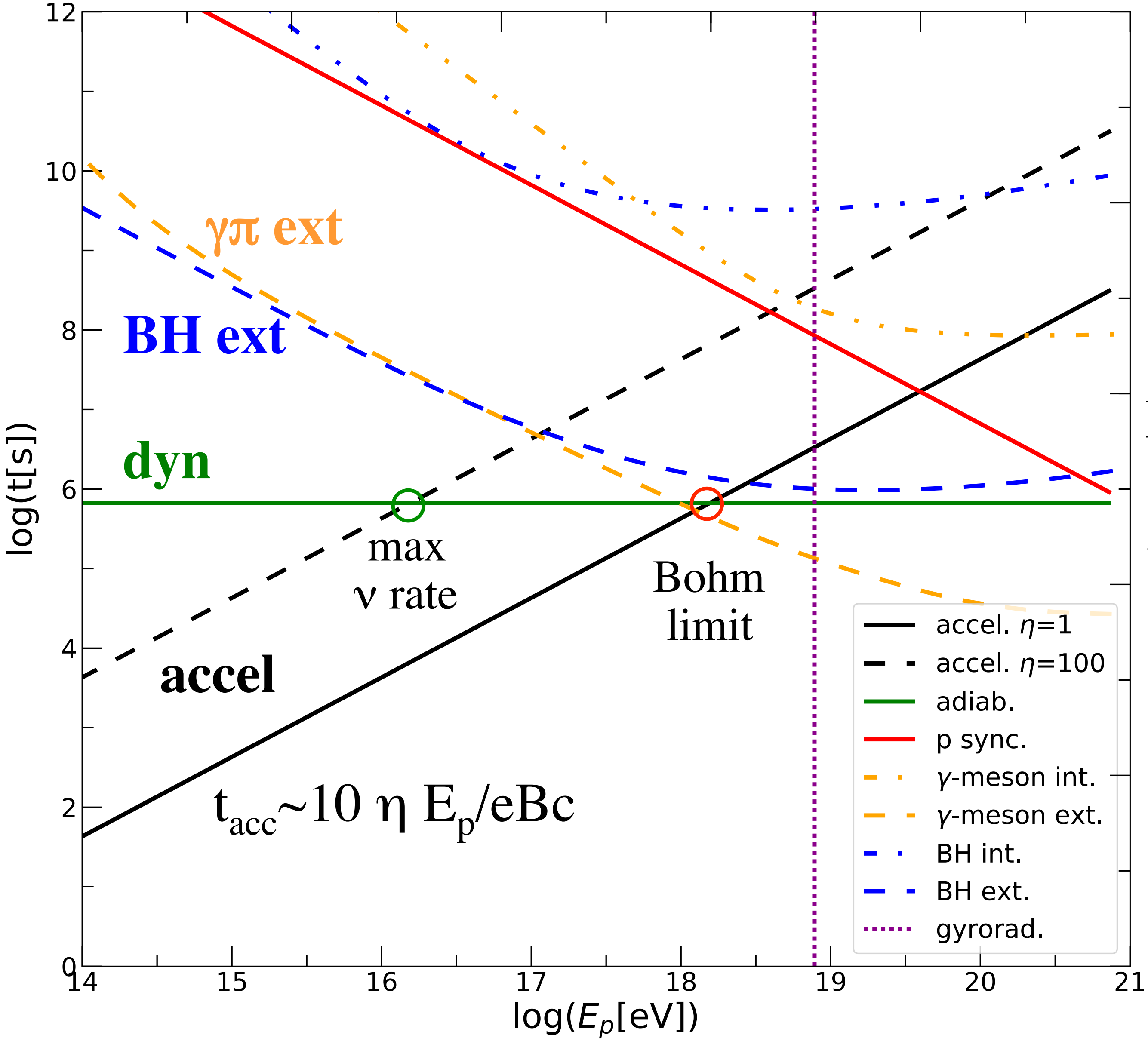
c.f. old (but prevalent) view





SI, Takahara 96

energy loss/accel. timescale comparison (comoving frame)

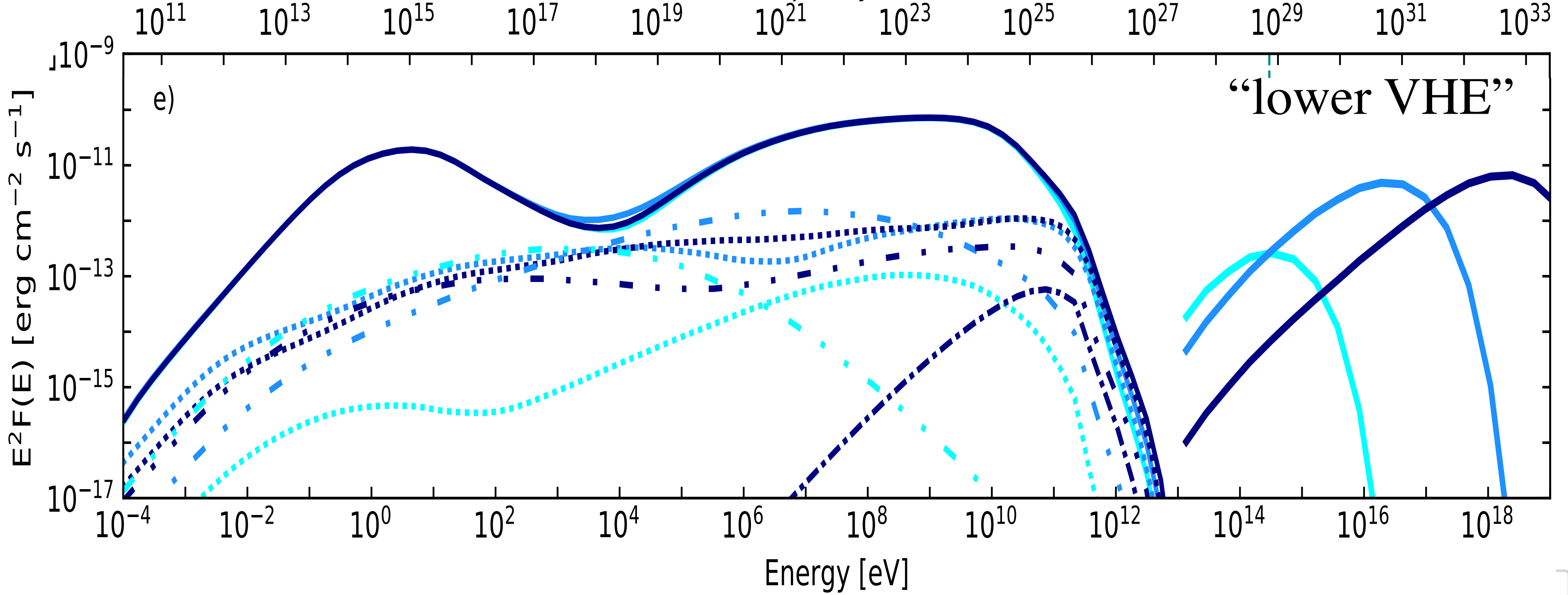


MAGIC 1807.04300

$E'_{pmax} \sim 10^{18}$ eV (comoving) achievable in principle -> UHECR?

implications: maximum proton energy

MAGIC 1807.04300



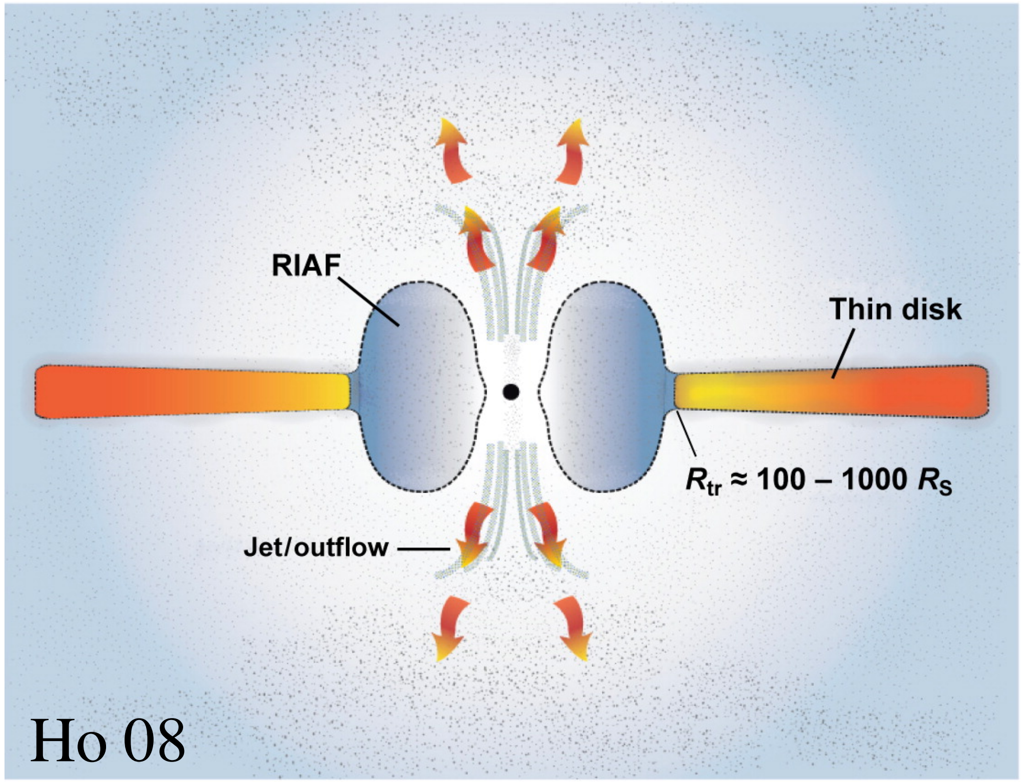
E_{pmax} dependence

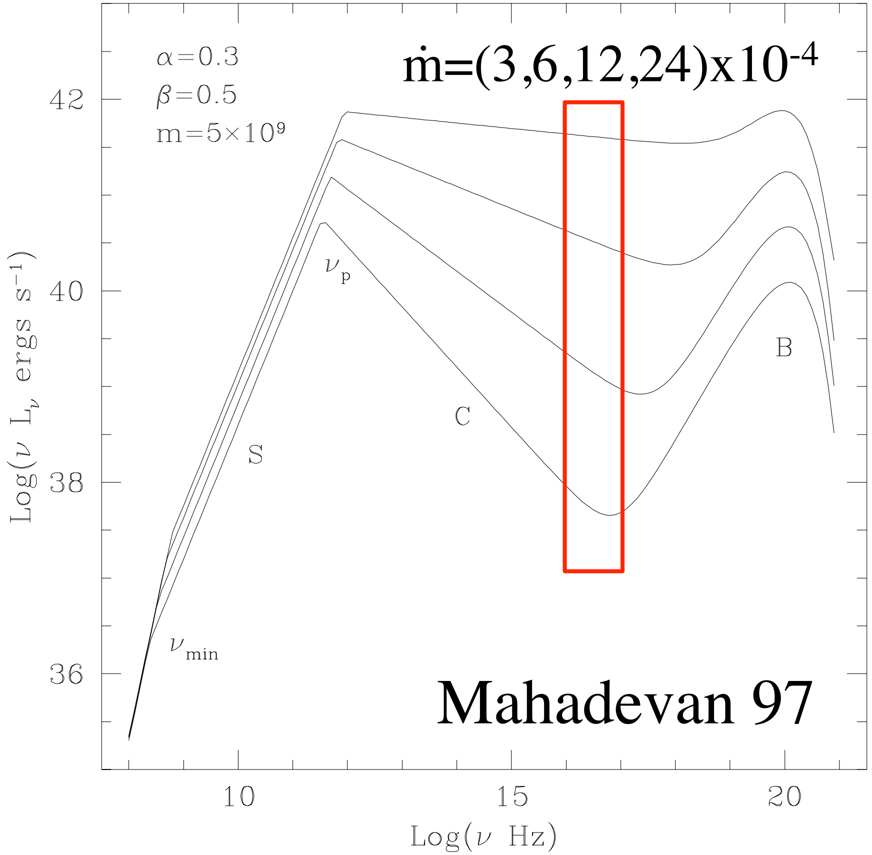
- maximum ν yield for E'_{pmax}~10^{16} eV (comoving)
- E'_{pmax}~10^{16}-10^{18} eV: higher $E_{\nu pk}$, lower L_p from X-ray limits
  lower EHE ν rate, but not too low to be ruled out (>~0.01/0.5yr)
  (simultaneously PS ν not too high, ~<0.05/0.5yr)
  -> unclear whether UHECR accelerator or not

3c. pγ scenarios with external photons from RIAFs

radiatively inefficient accretion flows

- expected at low accretion rates ($\dot{m}=\dot{M}/\dot{M}_{Edd}$~<0.01), inferred for SMBHs hosting BL Lacs
- radiatively inefficient -> hot, geometrically thick, optically thin <-> standard accretion disk for high $\dot{m}$
- broadband spectrum from radio to X-rays
- strong dependence of UV-soft X intensity on $\dot{m}$

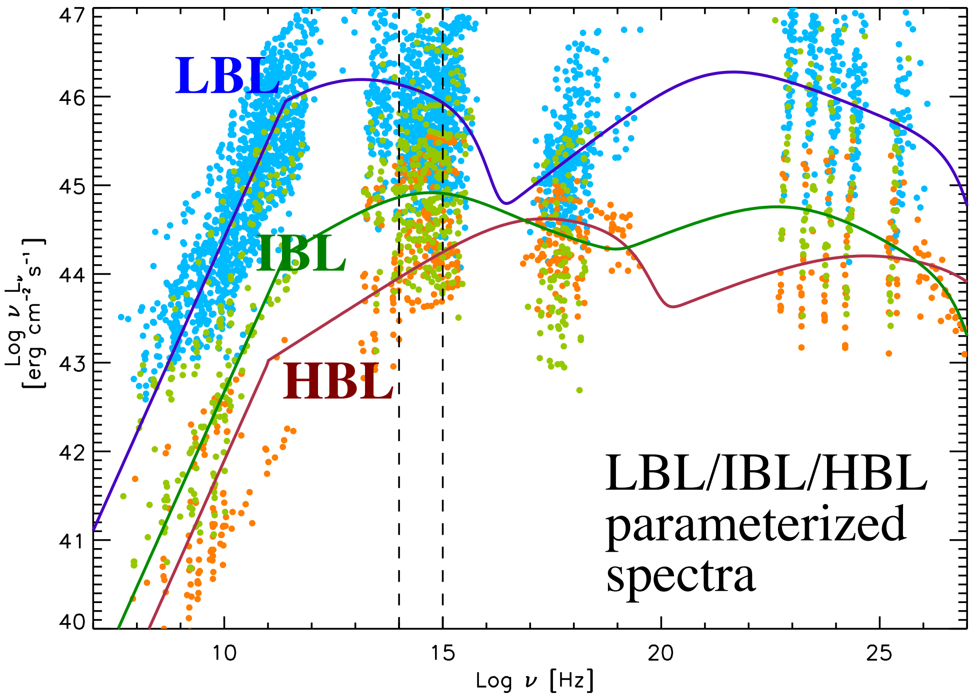


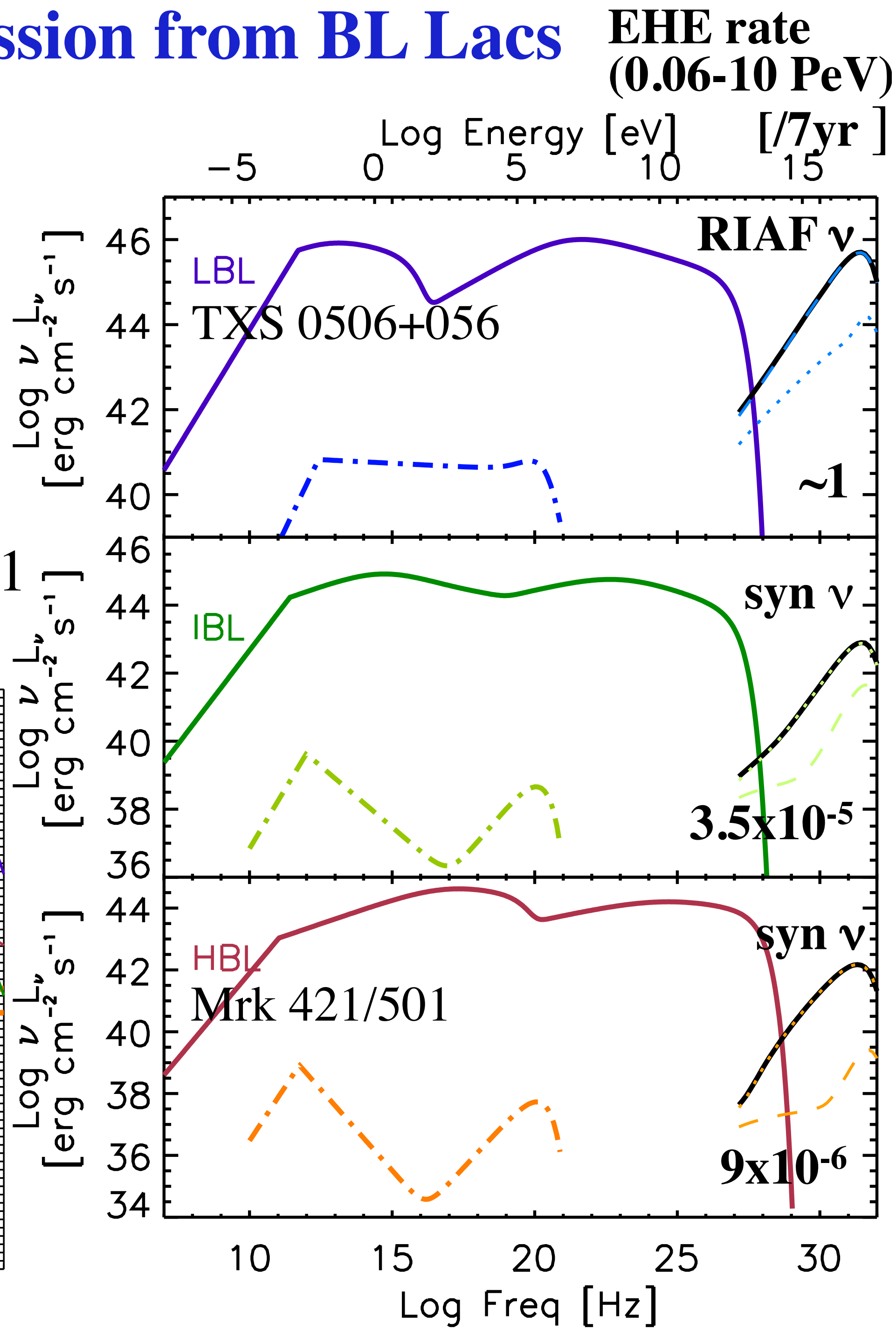


RIAF model for neutrino emission from BL Lacs

Righi, Tavecchio, SI 1807.10506

- plausible relation between P_{jet} and $\dot{m}$ of central SMBH among LBL/IBL/HBL subclasses
- RIAF model spectra well-defined
- effective pγ target only for LBLs
  -> potential explanation for why TXS 0506+056, not Mrk 421/501
- hadronic γ modeling to be done

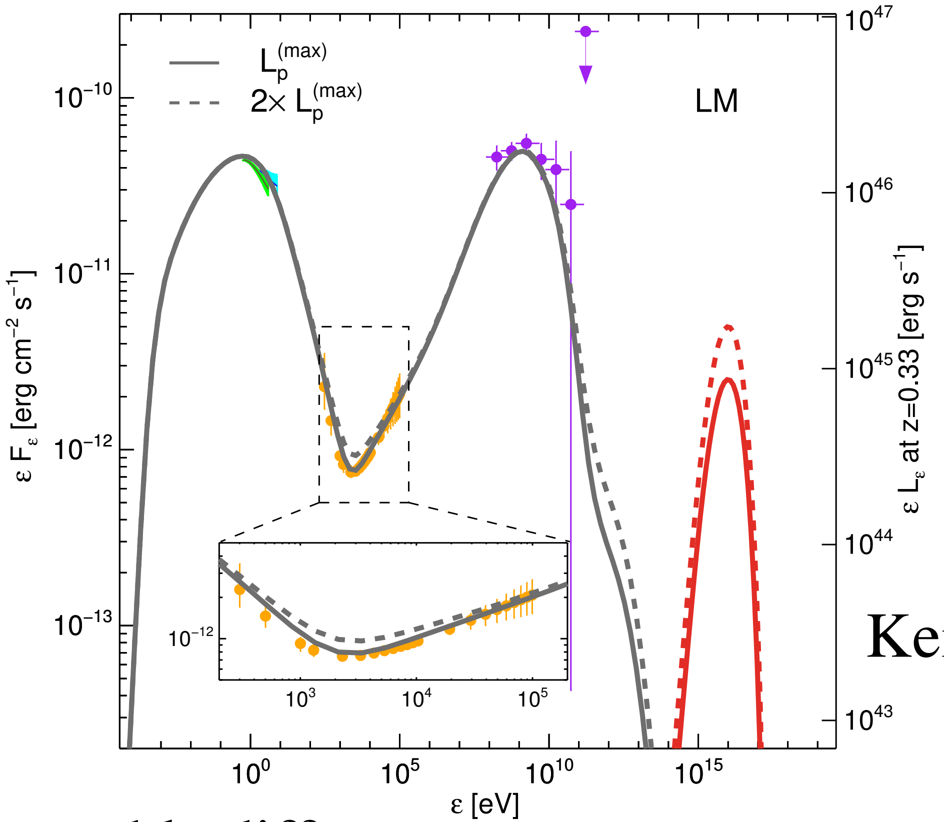


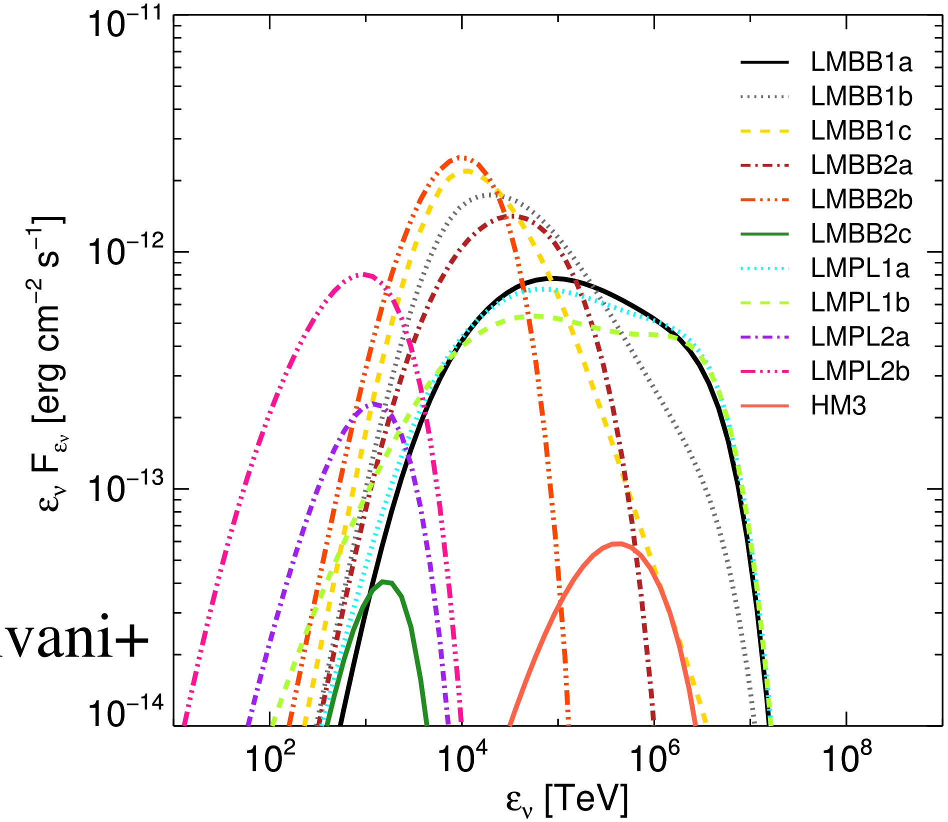


3d. brief comparison with other work

Keivani, Murase, Petropoulou+ 1807.04537

- generic internal/external photon scenarios
- hadronic-dominant scenarios difficult, likely leptonic-dominant
- X-ray cascade constraints, γ-ray absorption constraints crucial
- EHE rate <0.01/yr -> 1-zone models severely constrained if not ruled out





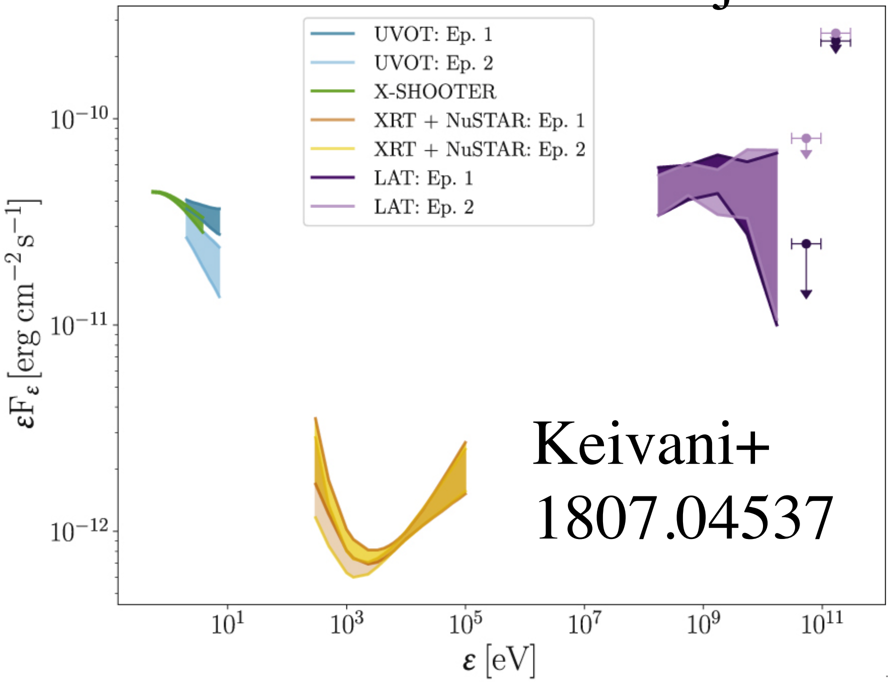
Keivani+

<u>notable differences</u>

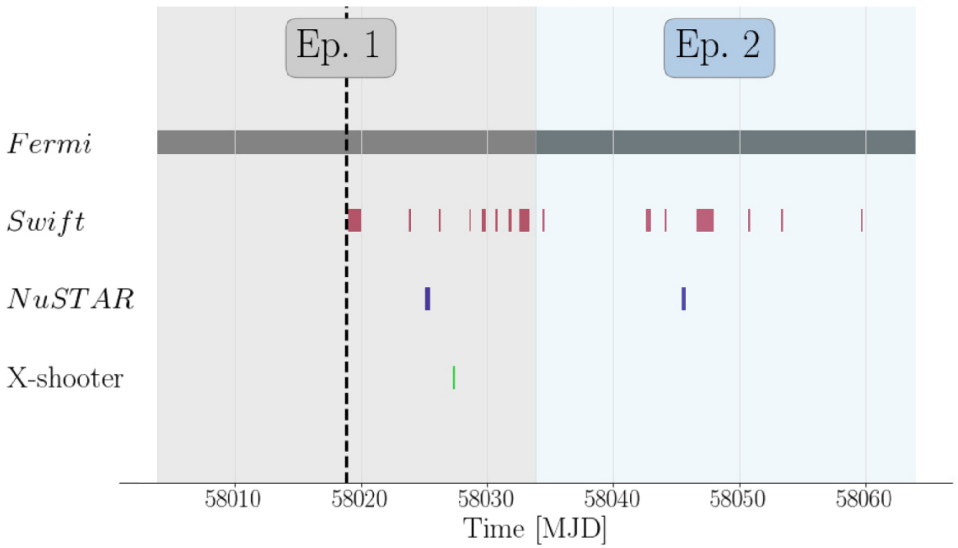
- choice of data set: different X-ray + UV, no VHE
- degree of optimism (~personal preference)

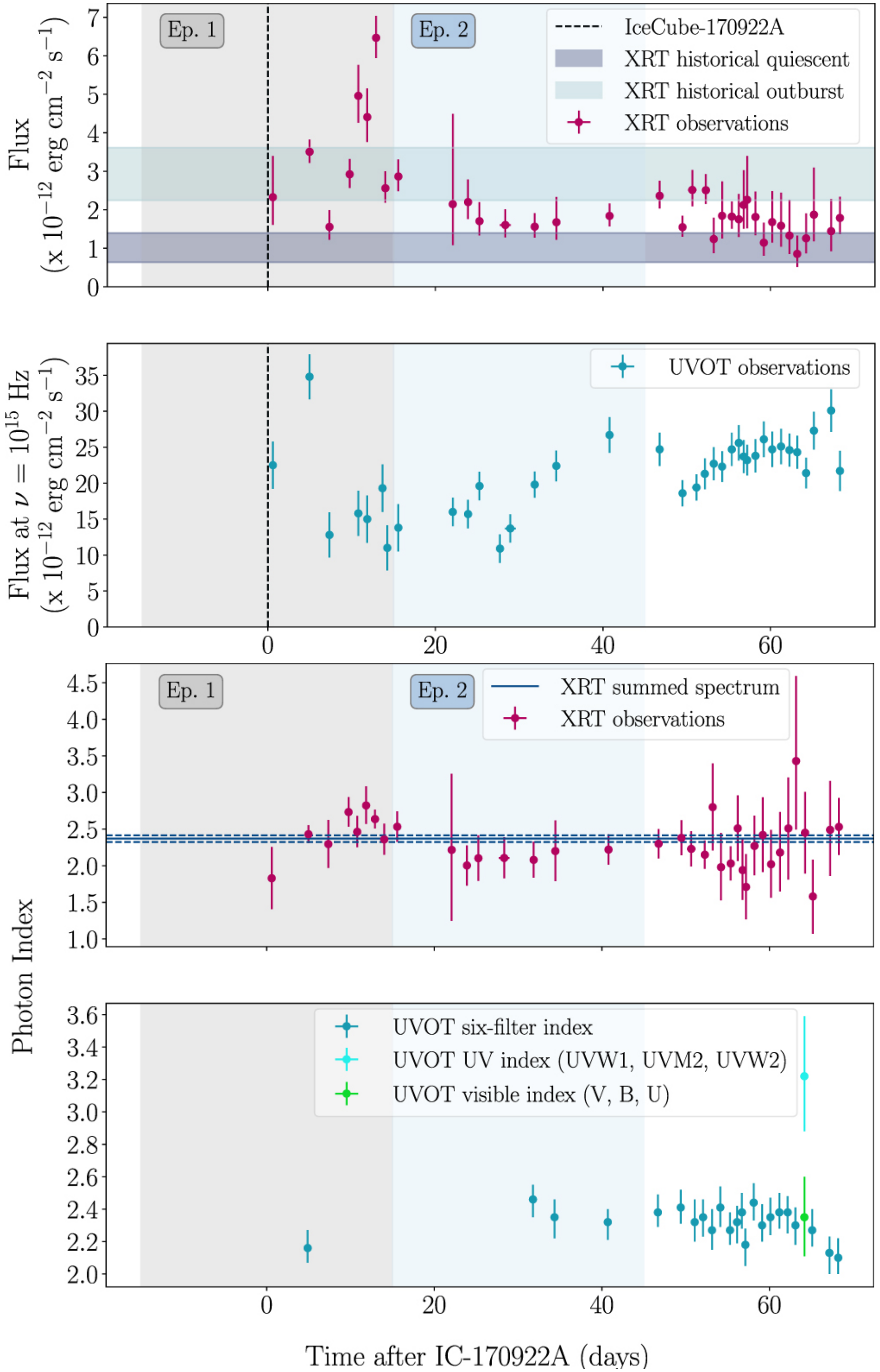
X-ray (+UV) data

with limited time coverage, exact choice of data set can be subjective…

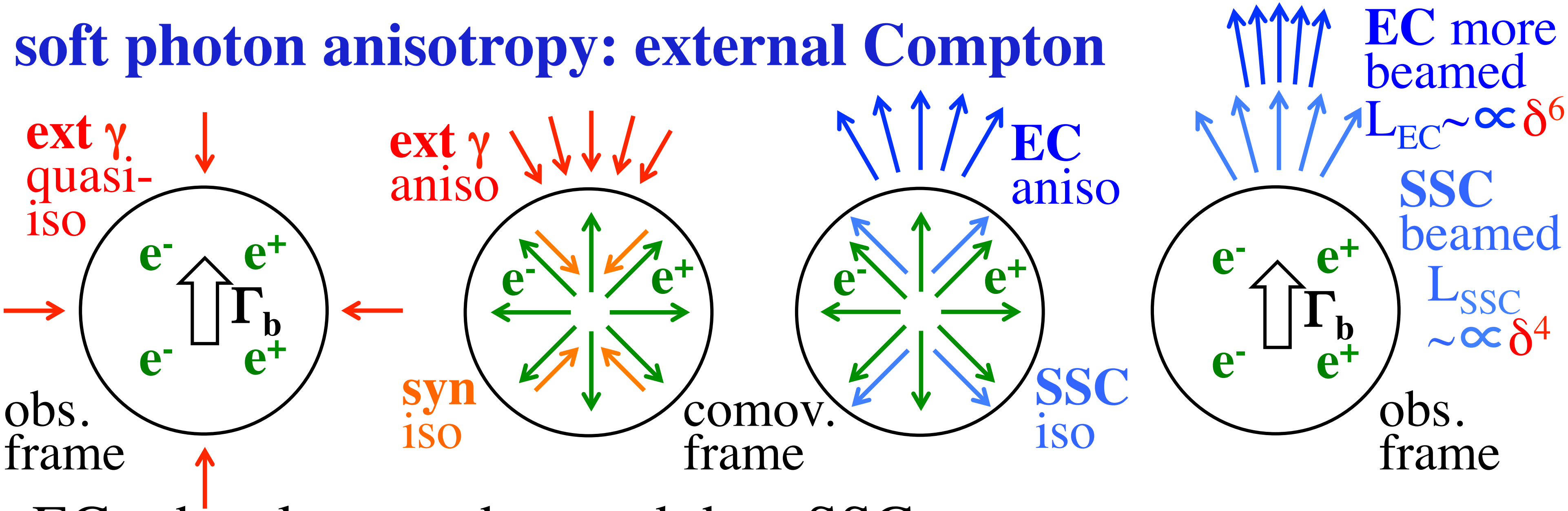


Keivani+ 1807.04537





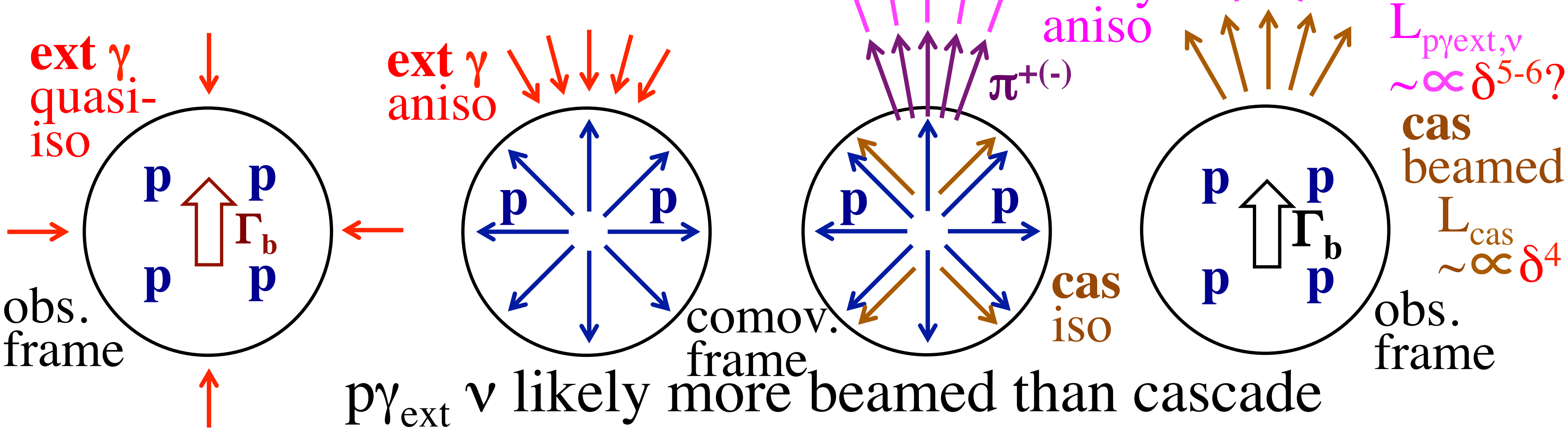
soft photon anisotropy: external Compton



EC robustly more beamed than SSC or syn. for quasi-isotropic external photons

Dermer 95

$p\gamma_{ext}$ neutrinos vs cascade?



$p\gamma_{ext}$ ν likely more beamed than cascade

summary electroweak observations/interpretation of TXS 0506+056

observations

- IceCube: detection of ~300 TeV neutrino, well localized
- Fermi-LAT: bright BL Lac TXS 0506+056 associated at 3σ CL
- MAGIC: <1 day variability, steep spectrum above ~100 GeV

interpretation via $p\gamma$ scenarios, one zone (e+p co-accelerated)

- internal target photons only: very high proton power required
- "external" photons from jet sheath: plausible & consistent
  - observed SED predominantly leptonic (sync.+external Compton)
  - hadronic subdominant, constrained by X-ray (+VHE)
    - GeV-TeV break consistent with $\gamma\gamma$ absorption entailed by $p\gamma$ production of ~300 TeV neutrino
    - proton max energy $\sim<10^{18}$ eV (comoving) possible in principle but not well constrained -> may or may not be UHECR accelerator
- external photons from RIAFs: also promising but only for LBLs
  -> potential explanation for why TXS 0506+056 and not HBLs
- addition of single neutrino to MWL SED provides crucial new insight

dawn(?) of electroweak astronomy

summary electroweak observations/interpretation of TXS 0506+056

questions

- relation to other blazars: why TXS 0506+056 and not HBLs, FSRQs?
- origin of 2014-2015 neutrino flare during low gamma-ray state (if real)
- contribution to diffuse flux, origin of dominant source(s)

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future

- more neutrino+EM observations necessary, especially X-rays, VHE
- more comprehensive modeling
- the game has just begun, further exciting times ahead!

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γ+ν+CRs: grand unified astronomy
γ+ν+CRs+GW: astronomy of everything